Formation of slowly rotating early-type galaxies via major mergers: a Resolution Study

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ABSTRACT
We study resolution effects in numerical simulations of gas-rich and gas-poor major mergers, and show that the formation of slowly-rotating elliptical galaxies often requires a resolution that is beyond the present-day standards to be properly modelled. Our sample of equal-mass merger models encompasses various masses and spatial resolutions, ranging from about 200 pc and 10^5 particles per component (stars, gas and dark matter), i.e. a gas mass resolution of $\sim 10^5$ M$_\odot$, typical of some recently published major merger simulations, to up to 32 pc and $\sim 10^3$ M$_\odot$ in simulations using 2.4×10^7 collisionless particles and 1.2×10^7 gas particles, among the highest resolutions reached so far for gas-rich major merger of massive disc galaxies. We find that the formation of fast-rotating early-type galaxies, that are flattened by a significant residual rotation, is overall correctly reproduced at all such resolutions. However, the formation of slow-rotating early-type galaxies, which have a low residual angular momentum and are supported mostly by anisotropic velocity dispersions, is strongly resolution-dependent. The evacuation of angular momentum from the main stellar body is largely missed at standard resolution, and systems that should be slow rotators are then found to be fast rotators. The effect is most important for gas-rich mergers, but is also witnessed in mergers with an absent or modest gas component (0–10% in mass). The effect is robust with respect to our initial conditions and interaction orbits, and originates in the physical treatment of the relaxation process during the coalescence of the galaxies. Our findings show that a high-enough resolution is required to accurately model the global properties of merger remnants and the evolution of their angular momentum. The role of gas-rich mergers of spiral galaxies in the formation of slow-rotating ellipticals may therefore have been underestimated. Moreover, the effect of gas in a galaxy merger is not limited to helping the survival/rebuilding of rotating disc components: at high resolution, gas actively participates in the relaxation process and the formation of slowly-rotating stellar systems.

Key words: galaxies: formation – galaxies: elliptical and lenticulars, cD – galaxies: interactions – galaxies: kinematics and dynamics – methods: N-body simulations

1 INTRODUCTION

Numerical simulations have been intensively used for more than two decades to study the properties of the remnants of galaxy mergers and the role of hierarchical merging in the formation of elliptical-like early-type galaxies (Hernquist & Barnes 1991; Barnes 1992; Mihos et al. 1995). With the increasing resolution and large statistical samples (e.g., Naab & Burkert 2003; Bournaud, Jog & Combes 2005; di Matteo et al. 2007, 2008), modern work tends to quantify in details the properties of major and minor merger remnants, and accurate comparisons with observed properties of early-type galaxies can now be envisioned (e.g., Burkert et al. 2008).

A general concern, though, is that the impact of the spatial and the mass resolutions on the detailed properties of the systems under scrutiny remains largely overlooked, and whether or not simulations of mergers have converged with today's typical resolution remains unexplored. Obviously, increasing resolution enables simulations to directly resolve cold gas clouds and clustered star formation (e.g. Bournaud, Duc & Emsellem 2008; Kim, Wise & Abel 2009), but whether these additional small-scale ingredients can significantly impact the global, large-scale morphology and kinematics of merger remnants has not been studied in detail. In cosmological simulations, an increase in resolution (i.e., an increase in the number of particles and/or decrease of the softening length) can affect the baryonic density and circular velocity profiles of individual galaxies in a halo (Naab et al. 2007). Navarro et al. (2008) also studied numerical convergence via a suite of ΛCDM simulations and confirmed that the halo mass distributions were better described by Einasto profiles that are not, stricly speaking, universal.

While many resolution studies have been made in cosmological simulations, few have focused on galaxy merger simulations. Cox et al. (2006), Hopkins et al. (2008) and e.g., di Matteo et al. (2008) included some checks of the effect of resolution on the star-formation activity of on-going mergers. But a resolution study aimed at examining the detailed morphology and kinematics of relaxed merger remnants (i.e., galaxies which tend to be roughly S0- or elliptical-like) has not yet been conducted.

Models of galaxy mergers have reached particularly high resolution with the work of Wetzstein, Naab & Burkert (2007) (70 pc softening with 4×10^6 particles in total – but only 45,000 for the gas component), Li, Mac Low & Klessen (2004) (10 to 100 pc and 5×10^5 gas particles per galaxy), Naab et al. (2007) (8×10^6 particles with a 125 pc resolution in a cosmological re-simulation of an individual galaxy halo). The highest resolution for gas-rich mergers have been achieved recently by Bournaud, Duc & Emsellem (2008) for mergers of bright spiral galaxies, with a total of 3.6×10^7 particles including more than 10^7 gas particle, and a 32 pc softening size, and Kim, Wise & Abel (2009) with a spatial resolution of 3.8 pc and a mass resolution of 2×10^3 M$_\odot$ (for dwarf or low-mass spirals, though). But such high-resolution studies have focused on small-scale gas physics and structure formation, without studying the impact of high resolution on the global properties of the elliptical-like galaxies formed in major mergers.

Large samples of simulations of idealized galaxy mergers remain typically limited to softening lengths of about 100–300 pc, and $\sim 10^5$ particles per galaxy (see samples in Naab & Burkert 2003; Naab, Jesseit & Burkert 2006; Bournaud, Combes & Jog 2004; Bournaud, Jog & Combes 2005; Cox et al. 2006, 2008; di Matteo et al. 2007). Whether or not the relatively limited numerical resolution used in such studies affects the global properties of

merger remnants is still a largely open question: for instance, the detailed comparison of major merger remnants with the observed anisotropy-flattening relation by Burkert et al. (2008) relies on simulations with 2×10^4 gas particles per galaxy, a gas particle mass $\sim 3 \times 10^5$ M$_\odot$, and a spatial resolution (softening) of about 200 pc.

Within the context of the ATLAS^{3D} project (http://purl.org/atlas3d), an extensive set of numerical simulations is being conducted to support the multi-wavelength survey of a complete sample of early-type galaxies within the local (40Mpc) volume, in terms of various formation mechanisms of early-type galaxies: binary mergers, multiple mergers, disc instabilities, etc. An ambitious series of simulations of mergers are being specifically performed and analysed for this purpose (Bois et al. in preparation). To properly interpret the results from these simulation efforts, as well as to understand the robustness of existing and past studies of galaxy mergers, we first probe the effect of spatial and mass resolutions on the global structure of binary disc merger remnants.

In this paper, we study the effect of numerical resolution on the global morphology and the kinematics of the simulated remnants of binary, equal-mass major mergers. We wish to examine resolutions ranging from the typical resolutions used in recent, large simulations samples, to some of the highest merger simulations ever performed. We study both dry (collisionless) and wet (gas-rich) mergers of disc galaxies. The modeled interaction orbits lead to the formation of both fast rotators, i.e., early-type galaxies flattened by significant rotational support, and slow rotators, i.e., early-type galaxies with low residual rotation, supported (and flattened) by (anisotropic) velocity dispersions, following the classification detailed in Emsellem et al. (2007, see also Section 2). We find that the formation of fast rotators is overall correctly reproduced with numerical simulations at modest resolutions. In contrast, the formation of slow-rotating systems is correctly reproduced only at high resolution (Section 3), above the resolution of most of the recently published merger simulations. The influence of gas on the structure of merger remnants, compared to dry mergers, also differs at high resolution, and is not limited to easing the survival and/or rebuilding of rotating disc components. In Section 4, we further examine the origin of this observed resolution effect in the formation of slow-rotating systems. We show that it is not an artefact from different initial conditions or interaction orbits, but that the physical treatment of the merging process is actually biased when the resolution is too low. The effect of the resolution has been tested on other simulations producing slow rotators and we find that it is a systematic one (Section 5). We summarize our results, discuss the required resolution for accurate studies and the general implications for the formation of elliptical galaxies in Section 6.

2 SIMULATIONS AND ANALYSIS

2.1 Method

2.1.1 Code

We use the particle-mesh code described in Bournaud, Duc & Emsellem (2008) and references therein.

This code uses a Cartesian grid on which the particles are meshed with a "Cloud-In-Cell" interpolation. The gravitational potential is computed with an FFT-based Poisson solver and particle motions are integrated with a leap-frog algorithm, and a time-step of 0.5 Myr.

Interstellar gas dynamics is modelled with the sticky-particle

scheme with elasticity parameters $\beta_t = \beta_r = 0.6$. This scheme neglects the temperature and thermal pressure of the gas, assuming it is dominated by its turbulent pressure, which is the case for the star-forming interstellar medium at the scales that are studied here (Elmegreen & Scalo 2004; Burkert 2006). The velocity dispersion of the particles model the turbulence, and their mutual collisions are inelastic to ensure that the turbulence dissipates over about a vertical crossing time (Mac Low 1999).

The star formation rate is computed using a Schmidt-Kennicutt law: it is then proportional to the gas density in each cell to the exponent 1.5. Gas particles are converted to star particles with a corresponding rate in each cell. Energy feedback from supernovae is accounted for with the scheme proposed by Mihos & Hernquist (1994). Each stellar particle formed has a number of supernovae computed from the fraction of stars above 8 M$_\odot$ in a Miller-Scalo IMF. A fraction ϵ of the 10^{51} erg energy of each supernova is released in the form of radial velocity kicks applied to gas particles within the closest cells. We use $\epsilon = 2 \times 10^{-4}$, as Mihos & Hernquist (1994) suggest that realistic values lie around 10^{-4} and less than 10^{-3}.

2.1.2 Set-up for initial disc galaxies

The baryonic mass of our model galaxies is 10^{11} M$_\odot$. In dry merger simulations, this mass is purely stellar. In wet merger simulations, 80% of this mass is stellar, and 20% is gaseous. The initial gas and stellar discs are Toomre discs, with a scale-length of 4 kpc and a truncation radius of 10 kpc for the stars, resp. 8 kpc and 20 kpc for the gas. 20% of the stars are in a spherical bulge, modeled with a Hernquist (1990) profile with a 700 pc scale-length. The dark matter halo is modelled with a Burkert profile (Burkert 1995), a 7 kpc scale-length and a truncation radius of 70 kpc, inside which the dark matter mass is 3×10^{11} M$_\odot$.

The two "progenitor" disc galaxies in each simulation are identical, the total mass of the remnant will be 2×10^{11} M$_\odot$ which is consistent with the *slow-rotator* mass range observed in the ATLAS^{3D} sample (Emsellem et al. 2007).

2.1.3 Orbits

We have used two interacting orbits, for each kind of merger (dry and wet) and each resolution level. None corresponds to a very specific and unlikely configuration like coplanar discs, or polar orbits.

The first orbit is called "fast" because it forms fast-rotating early-type galaxies. The velocity at an infinite distance is 170 km s^{-1} and the pericentre distance is 30 kpc. This orbit is prograde with respect to the first progenitor disc, with an inclination of the orbital plane w.r.t. the disc plane of 25°. The orbit is retrograde w.r.t. the other progenitor disc, with an inclination of 45°.

The second orbit is called "slow" because it forms slow-rotating early-type galaxies (at least at high-enough resolution). The velocity at an infinite distance is 140 km s^{-1} and the pericentre distance is 25 kpc. This orbit is prograde with respect to the first progenitor disc, with an inclination of the orbital plane w.r.t. the disc plane of 45°. The orbit is retrograde w.r.t. the other progenitor disc, with an inclination of 25°.

These orbits as well as those used in additional tests (Section 5) have a total energy $E > 0$ or $E \simeq 0$, corresponding to initially unbound galaxy pairs. Such orbits are representative of the most common mergers in Λ-CDM cosmology (Khochfar & Burkert 2006).

Label	Softening Length	Particles / component	Total particles
standard	180 pc	10^5	6×10^5
high	80 pc	10^6	6×10^6
very high	32 pc	6×10^6	3.6×10^7

Table 1. Label for the resolution, softening length, number of particles per component (stars, gas and dark matter) and total number of particles in the simulation for the three resolutions.
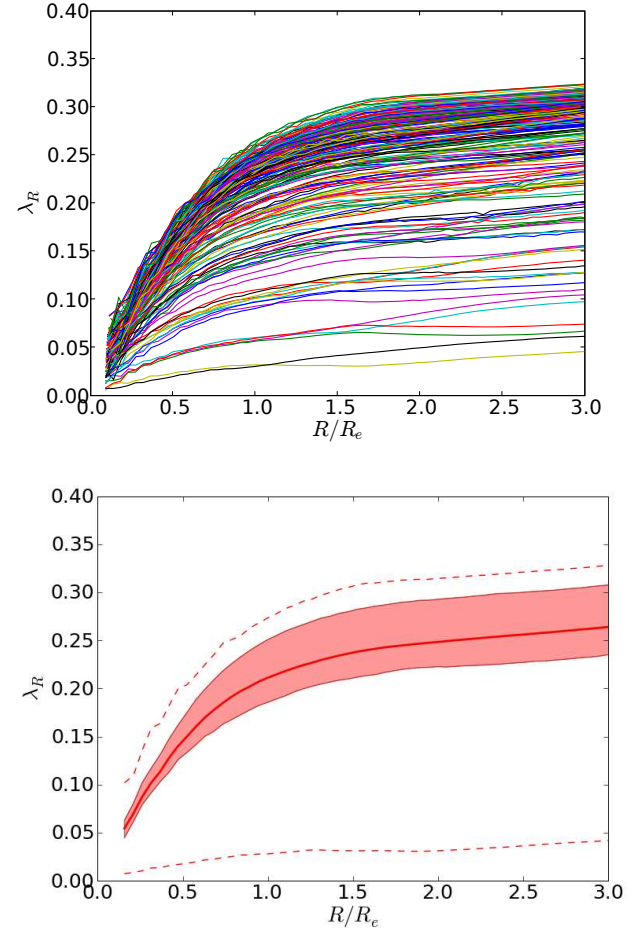

Figure 1. **Top panel** λ_R profiles of all 200 projections for the *Dry-Fast-very High* simulation. The profiles are plotted as a function of R/R_e (radius normalised by the effective radius for each projection). Each line here corresponds to a given projection. **Bottom panel** Corresponding median (thick solid line), quartiles (thin solid line) and maximal and minimal values (dashed lines) at each radius. This representation of the results illustrates the fact that all values are between the dashed lines, and 50% of the projections are in the filled area. Each line plotted in this panel does not correspond to one specific projection: the median or quartiles are derived for different projections at each radius.

2.1.4 Standard-, high- and very high-resolutions

Wet and dry mergers have been simulated for each orbit at three resolution levels. The detail for these resolutions are indicated in Table 1. The *very high* resolution arguably corresponds to the highest resolution simulation of a wet major merger performed so far (see Bournaud, Duc & Emsellem 2008).

We will label each simulation with the following nomenclature:

- the first item indicates a *Wet* or *Dry* merger, i.e. gas-rich or gas-free progenitors.
- the second item specifies the chosen orbit: the one producing *fast-rotators* or *slow-rotators* (at least at high-enough resolution).
- the last item indicates the resolution level: *standard*, *high* or *very high*.

For instance, the *wet-fast-high* simulation refers to the high-resolution models of a wet merger on the orbit producing a fast-rotating early-type galaxy.

2.2 Analysis of the relaxed merger remnants

We analyse the merger remnants after 1.2 Gyr in the simulation, which is 800-900 Myr after the first pericentre passage, and 600-700 Myr after the central coalescence. The remnants are thus relaxed when the analysis is performed. Tidal debris can still be orbiting around the merger remnant, but the bulk of the stellar mass in the central body does not show significant evolution. Analysis performed at earlier and later instants did not show significant variations, so spurious effects related to time evolution should not affect the comparison of the three different resolution levels.

2.2.1 Projected maps

Intrinsic and apparent properties of the merger remnant (e.g the apparent ellipticity) are directly linked with its orbital structure (Jesseit, Naab & Burkert 2005). To probe the relaxed merger remnants, we have therefore built projected maps of the stellar mass density, line-of-sight velocity and velocity dispersion fields. Two-dimensional maps are useful to reveal the wealth of photometric or kinematic structures associated with a galaxy merger, e.g., globular clusters or kinematic misalignments (see Bendo & Barnes 2000; Jesseit et al. 2007).

The projected maps cover a 16×16 kpc^2 field of view around the density peak of each system: our analysis is conducted up to a limit of three effective radii R_e, which encloses most of the baryonic mass of early-type galaxies, and the typical effective radius of our merger remnants is 2.5 kpc. Each projection was computed on a 100×100 pixel grid. The pixel size is 160×160 pc^2, which approximatively corresponds to the size of the softening length of our *standard-res* simulations, and is kept fixed for all resolutions.

To obtain statistically significant results, we have built such maps and performed the subsequent analysis with 200 isotropically distributed viewing angles (i.e., 200 different line-of-sights). In this way, we do not characterize and compare the merger remnant under a particular projection, but their global, statistical properties. As an example, Figure 1 shows the effect of the projections on the radial λ_R profiles for one simulation. Among these 200 profiles, the lowest (near zero) and highest values correspond respectively to the merger remnant seen nearly face-on (i.e., the lowest apparent ellipticity) or nearly edge-on (i.e., the highest apparent ellipticity). Our choice of 200 projections ensures that neighbouring projections are separated only by 10 degrees in any direction, so that intermediate viewing angles would not show significant differences.

2.2.2 Physical parameters

Our analysis is based on a few simple morphological and kinematic parameters – a choice mainly motivated by the fact that these parameters are often being used as standards in studies of nearby elliptical galaxies.

The morphological parameters pertains to the photometry : we measure the ellipticity ϵ (defined as $1 - b/a$, where a and b are the semi major- and minor-axes, respectively) and a_4/a which is the fourth (cosine) Fourier coefficient of the deviation of isophotes from a perfect ellipse ($a_4/a > 0$ and $a_4/a < 0$ correspond to discy, and boxy isophotes, respectively). These two parameters are computed using the Kinemetry software tool[1] which can be used to perform standard ellipse-fitting of galaxy images, as well as to study galaxy kinematics (Krajnović et al. 2006). For the kinematic analysis, apart from the first two velocity moments (velocity and velocity dispersion), we use the λ_R parameter, a robust proxy for the baryonic projected angular momentum, as defined in Emsellem et al. (2007, hereafter E07) :

$$\lambda_R \equiv \frac{\langle R \,|V| \rangle}{\langle R \sqrt{V^2 + \sigma^2} \rangle}$$

In E07, λ_R was used to reveal two families of early-type galaxies, the *slow-rotators* with $\lambda_R \leqslant 0.1$ and the *fast-rotators* with $\lambda_R > 0.1$ at one effective radius R_e. In a recent study, Jesseit et al. (2008) have simulated binary disc mergers to investigate the λ_R parameter: tests on their merger remnants reveal that λ_R is a good indicator of the true angular momentum content in early-type galaxies. As emphasised in E07, Cappellari et al (2007) and Krajnović et al. (2008), fast and slow rotators exhibit qualitatively and quantitatively different stellar kinematics. λ_R is thus an interesting parameter to probe, and should indicate whether or not the kinematics of the merger remnants are equally resolved at different resolutions.

For each above-mentioned parameter, we have computed the minimum, maximum, mean values, as well as the 1st and 3rd quartiles over all the projections at individual radii, to quantify the statistical distribution of these parameters in a simple way. An example is shown in Figure 1. Note that with this choice, the projection which minimises or maximises a parameter varies with radius.

3 EFFECT OF RESOLUTION ON THE FORMATION OF SLOW ROTATORS

In this Section, we briefly describe the properties of the simulated mergers with the three different resolutions. The complete set of analysis results can be found in Appendix A. We then focus the analysis on the simulations that show important differences, namely the cases producing slow rotators at high resolution.

3.1 Morphology and kinematics

Radial stellar density profiles are shown in Figure 2. We then show in Figure 3 the projected stellar density maps, of the relaxed merger remnants in all simulations, choosing the flattest and roundest projections as well as a projection representative of the mean ellipticity in each case. The corresponding line-of-sight stellar velocity fields are presented in Figure 4 for the same projections, the maps have been Voronoi binned (Cappellari & Copin 2003) to the same level of 15 particles minimum per bin. Further morphological or kinematics parameters are presented in Appendix A.

This analysis reveals various similarities or differences, depending on which merger is considered. The most noticeable results are as follows:
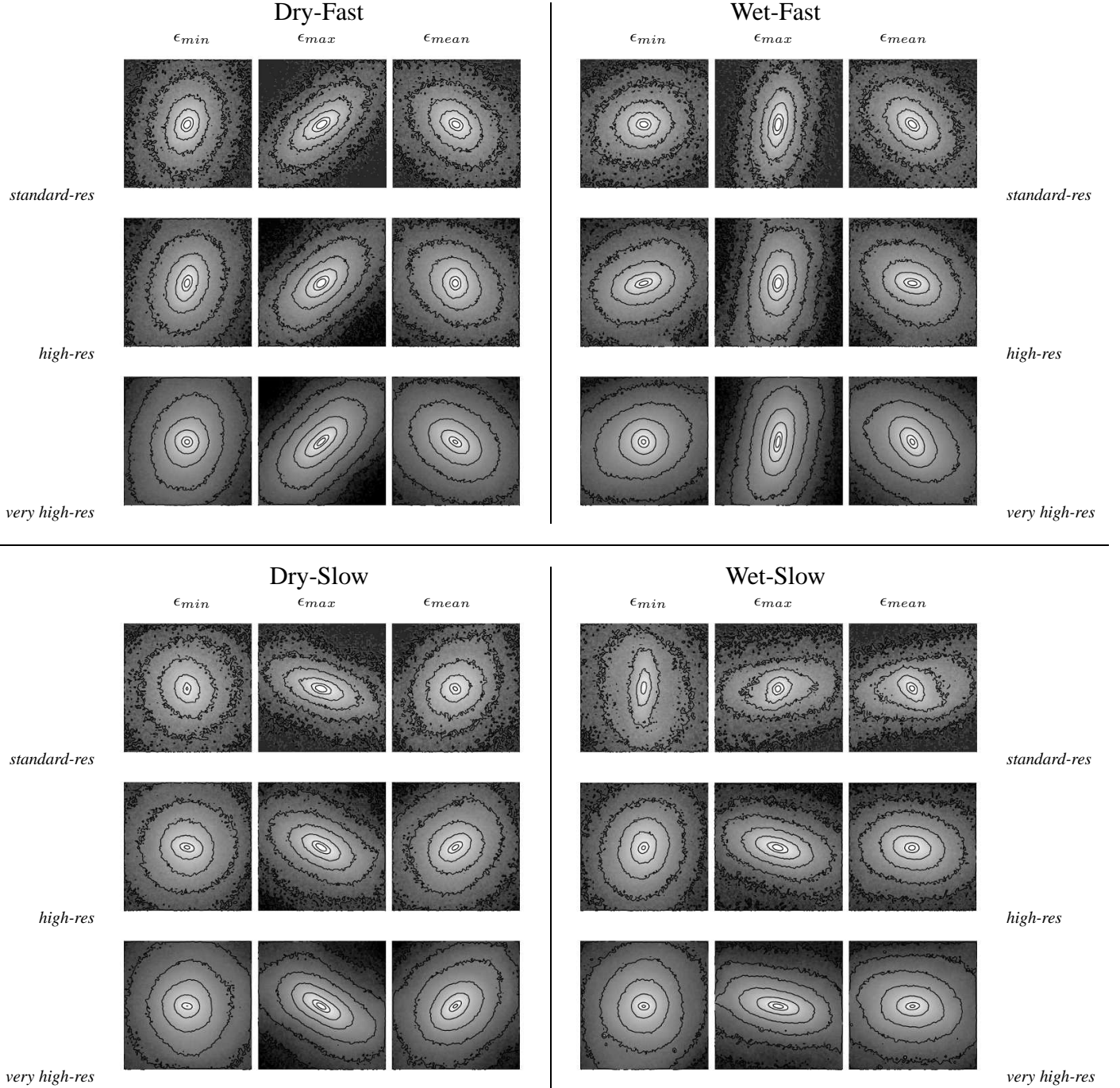
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[1] http://www-astro.physics.ox.ac.uk/ dxk/idl/

Figure 3. The twelve normalized projected surface density maps (in log), for the four sets of simulations at three different resolutions (labelled accordingly). The field of view is 16×16 kpc^2. For each simulation, the projections corresponding to the minimum, maximum and mean ellipticities are shown. The viewing angle of these projections are defined at *very high-res* and re-applied for the *standard-res* and *high-res* simulations: projections are thus established along the same line-of-sights for all resolutions. Luminosity contours are the same for all simulations and drawn with a spacing of 0.5 magnitude (except for the 2 inner contours with a step of 0.3). The effective radius is about at the edge of the 4th isophote for all simulations.

● *Mergers that produce fast-rotators at the highest resolution also result in fast rotating systems at the lower, standard resolution.* Overall, the apparent morphology for any projection of the Dry-Fast and Wet-Fast models is unaffected by the resolution (Figure 3). The velocity fields are also quite similar (Figure 4), with only minor misalignments between the apparent kinematic and photometric axes. Ellipticity and λ_R profiles, provided in Appendix A (see

Figure A1 and A2), confirm these similarities and that all these mergers remnants are fast rotators, with a rotational support that is largely independent from the numerical resolution.

● *Strong kinematic misalignments and kinematically decoupled cores (KDCs) are found only in slow-rotators, but really appear only at high resolution.* The Dry-Slow model has a KDC at standard resolution, but its amplitude is significantly lower than the one
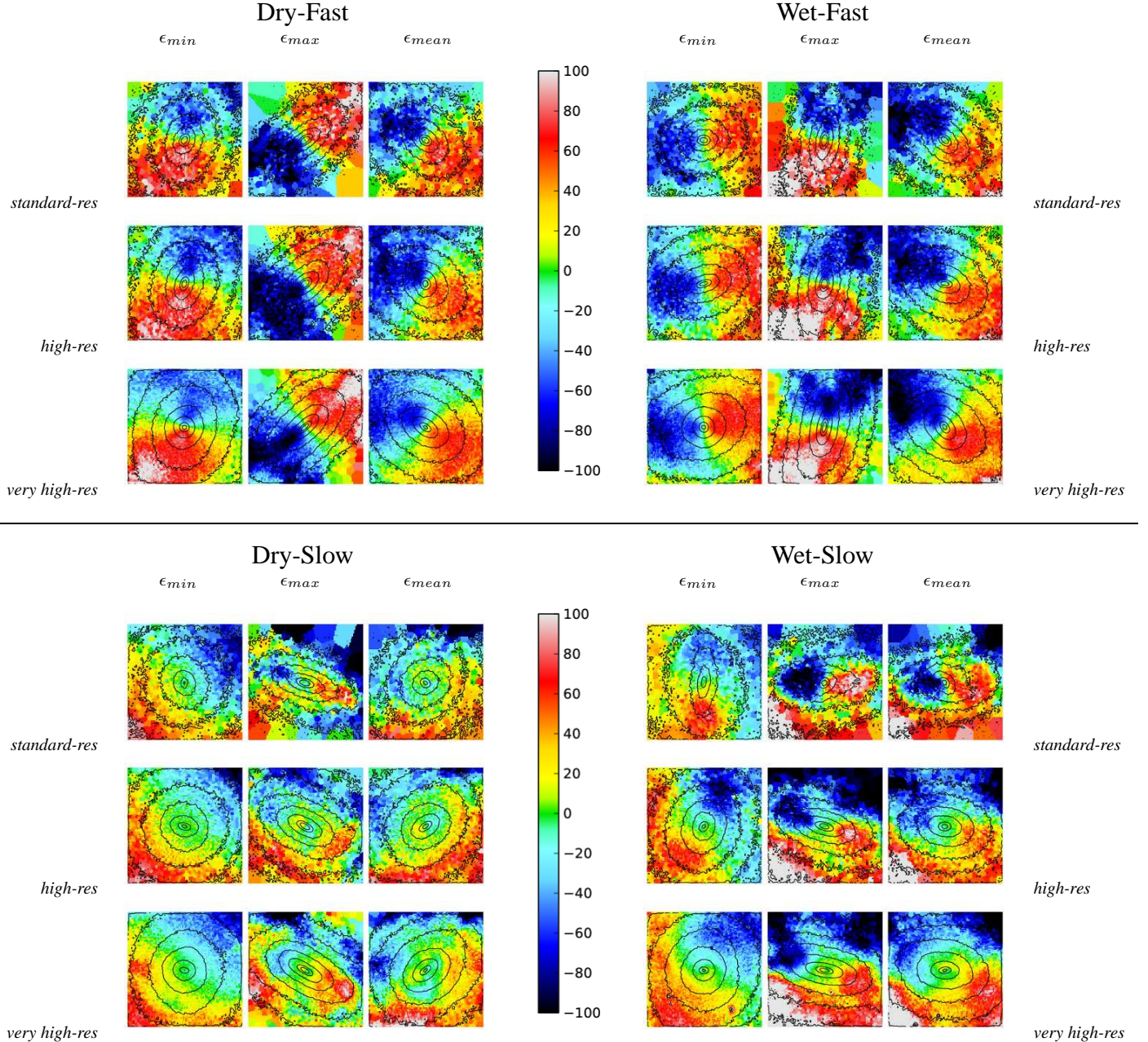
Figure 4. The twelve projected stellar velocity fields. The field of view is 16×16 kpc^2, projections and contours are the same as in Figure 3.

observed in the high and highest resolution models. The Wet-Slow model has a KDC only at high/very high resolution. Overall, kinematic misalignments increase at high resolution, as illustrated for instance by the flattest projections of the Wet-Slow case.

• *Morphological and kinematic differences are most important for mergers that produce slow-rotators at high resolution.* Striking morphological differences are seen in particular for the Wet-Slow case (Figures 2 and 3) and both the amplitude and the shape of the velocity field change with resolution for the Wet-Slow and Dry-Slow cases (Figure 4). For instance a rapidly rotating core is seen in the Wet-Slow merger remnant at standard-resolution, instead of a slow-rotating KDC at high- and very high-resolutions. The Dry-Slow remnant also shows up as a discy rotating system at standard resolution, in contrast with the observed remnant at higher resolu-

tions. We also note on Figure 2 that the stellar density profile is resolution-dependent in particular for the Wet-Slow case, with a much less concentrated merger remnant in the standard-resolution case (the mass within 5 kpc is about 25% lower than at high or very high resolutions).

3.2 Formation of *slow-rotators* at high resolution

We now focus on the detailed properties of the mergers for which the most important differences have been noticed, namely those producing slow rotators at the highest resolutions.
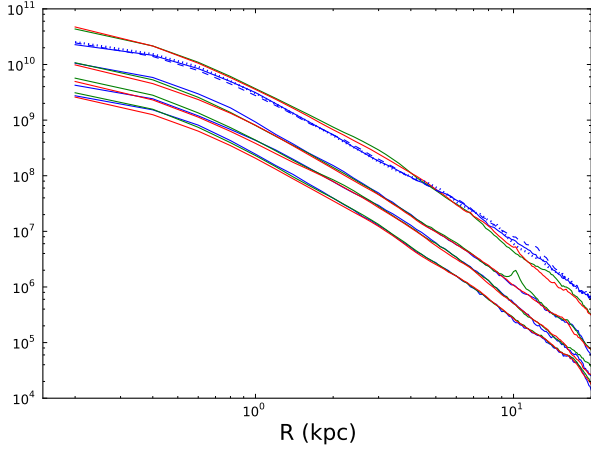
Figure 2. Stellar density profiles for the Wet-Slow, Wet-Fast, Dry-Slow and Dry-Fast simulations (from top to bottom, respectively). The density of the Wet-Fast, Dry-Slow and Dry-Fast cases has been divided by a factor of five, ten and twenty respectively to improve the readability of the plot. Red lines correspond to very high-resolution models, green lines to high-resolution cases, and blue lines to standard-resolution models. The two alternative realisations of the Wet-Slow-Standard simulation (see Sect. 4.3) are shown in dashed and dotted lines.

Figure 5. Projected density maps of the old stars, young stars, and gas in the wet-slow remnants; from top to bottom : *standard-*, *high-* and *very high-res* models, for the projection which minimises the ellipticity (ϵ_{min}) as in Figure 3. Old stars are those formed before the merger, young stars are formed during/after the merger. The field of view is 16×16 kpc^2, and the isocontours correspond to the projected old stellar component.
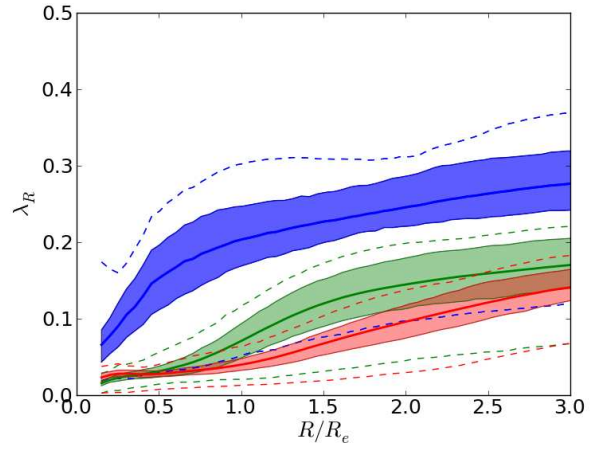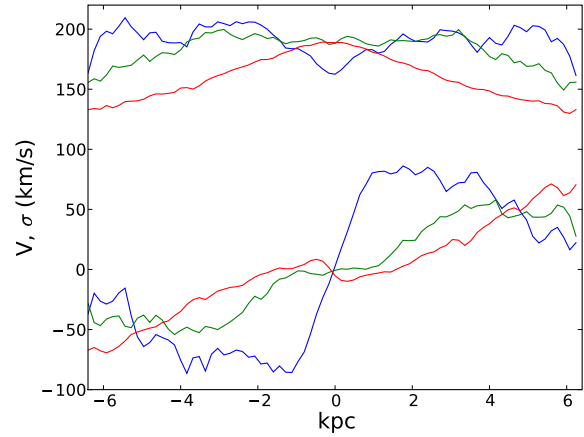

Figure 6. Top panel Radial line-of-sight velocity and velocity dispersion profiles (respectively bottom and top lines of the plot) for the mean ellipticity projection along the global kinematic position angle of the Wet-Slow simulation (right panels of Figure4). **Bottom panel** λ_R profiles as a function of R/R_e, the minimum, median, maximum and quartiles values are presented as in Figure 1. In both panels, the *standard-res* is represented in blue, the *high-res* in green, the *very high-res* in red.
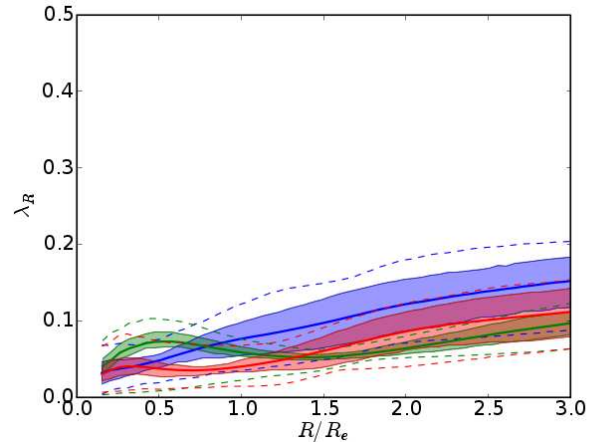

Figure 7. λ_R profiles as a function of R/R_e for the Dry-Slow simulation. The *standard-res* is represented in blue, the *high-res* in green, the *very high-res* in red.

3.2.1 Morphology and Kinematics

To better understand the differences seen in the morphology of the Wet-Slow simulations, we have examined the three included baryonic components of the merger remnants separately, namely the "old" stars formed before the beginning of the merger event, the "young" stars formed during/after the merger event, and the gas left over after the merger (see Figure 5). Within the central 10 kpc, the *standard-res* remnant exhibits a prominent bar, the inner distribution of the gas and young stars being driven by this tumbling structure with e.g., a ring-like structure at a radius of ~ 6 kpc. In the *high-res* and *very high-res*, the gaseous component and the young stars have a smoother distribution more closely following the overall old stellar distribution. In addition, many young star clusters are visible in the maps from the *very high-res*, a few in the *high-res*, and none in the *standard-res*. High spatial resolution of course allows to resolve the formation of stellar clusters (see also Bournaud, Duc & Emsellem 2008) but there is also a larger number of other young stellar substructures at increasing resolution, like filaments, tidal streams, and a compact nucleus (Figure 5).

The kinematic discrepancies discussed above in the velocity fields are quantified globally in the radial velocity and λ_R profiles (Figure 6). The *standard-res* displays significant rotation inside 3-4 kpc (up to ~ 85 km s^{-1}) and a decreasing rotation velocity at larger radii. There is a drop in the velocity dispersion in the central 2 kpc, and no sign of a KDC. This is in stark contrast with both the *high-res* and *very high-res* which overall show much lower rotational velocity support (below ~ 50 km s^{-1} and particularly low in the central 2 kpc), and a KDC in the central 1 kpc. Overall the *high-res* and *very high-res* have similar velocity rotation curves, apart from a more pronounced KDC signature in the *very high-res* (partly due to the KDC having a slightly different position angle in these two remnants).

The general discrepancies of the *standard-res* versus *high-res* and *very high-res* realisations are confirmed by the λ_R profiles (Figure 6). The merger remnant made at *standard-res* is clearly a *fast-rotator*. The *high-res* and *very high-res* are both classified as *slow-rotators* with respectively a maximum value of λ_R of 0.1 and 0.06 at one effective radius. The λ_R profile goes up somewhat more rapidly with radius in the *high-res* case than in the *very high-res*, but the difference remains of the order of the scatter between different projections of each case. The presence of a bar in the stellar component of the *standard-res* is likely a result of the significantly higher rotational support (see also Sect.5).

Beyond one R_e, the λ_R profiles of the *high-res* and *very high-res* are rising: there is less angular momentum in the center, which has been expelled outwards (see also E07). However, even at these large radii, the *slow-rotators* have less angular momentum than *fast-rotators* (see Figure 6). Observations conducted up to two or three R_e (Weijmans et al. 2009; Coccato et al. 2009) would bring additional constraints on the formation scenario of slow-rotating early-type galaxies.

3.2.2 Role of gas on the properties of merger remnants

The Dry-Slow simulations show smaller differences in the stellar density maps and velocity fields. They also exhibit smaller differences in their λ_R profiles (Figure 7). Nevertheless, the *standard-res* simulation is again a faster-rotator than the *high-res* and *very high-res* cases at 1, 2 or 3 effective radii. A KDC is also found only in the *high-res* and *very high-res* cases, associated to a peak of λ_R inside one effective radius.

A lower specific angular momentum in the main stellar body at higher resolution is not only found in Wet-Slow mergers, but also in Dry-slow mergers, the differences being still much more pronounced in the Wet case.

Gas plays an important role in shaping merger remnants (Naab, Jesseit & Burkert 2006; Robertson et al. 2006; Hopkins et al. 2009b) and it is interesting to compare the Wet-Slow and Dry-Slow merger remnants at fixed resolution, to better understand its specific impact (Figures 6 and 7):

- at *standard-res*, the wet merger remnant has a much higher rotational support than the dry case. This is consistent with the usually known effect of gas helping the survival of rotating stellar discs during major mergers, and/or re-building of discs after mergers (Robertson et al. 2006, Hopkins 2009b).
- at *high-res* and *very high-res*, the rotational support of the merger remnant is not increased when gas is present. The angular momentum, traced by λ_R, is actually lower by about 20% inside one effective radius in the *very high-res* wet case, compared to the corresponding dry merger.

It thus seems that the impact of gas on the global properties of major merger remnants is more complex than originally thought, and can even be weakened at high resolution. This suggests that the global dynamics of gas during the major merger or in a young merger remnant can be significantly affected by resolution. As seen in Figure 5, gas at *standard-res* largely lies in smooth structures and the formation of new stars during the merger proceeds in a relatively smooth way. At increased resolutions, thinner gas structures are resolved during the merger, which can result in clustered star formation and the formation of numerous young stellar structures, as observed in the final merger remnant in Figure 5.

3.3 Summary of the resolution tests

The resolution does not seem to significantly affect the morphology and kinematics of the mergers remnants that are fast rotators at high resolution: they are still fast rotators at lower resolution, with very similar morphological and kinematic properties. This contrasts with the fact that resolution has a major effect on the formation of slow-rotating systems. The systems that are slow rotators at high resolution rotate more rapidly when the resolution decreases, and can be observed as true fast rotators at *standard-res*. The effect is small in dry mergers, but is dramatic in our wet merger model. KDCs in these slowly-rotating systems are also significantly better resolved at high resolution. The role of gas in shaping merger remnants is found to vary with resolution: at low resolution, gas rebuilds rotating disc components, increasing the overall disciness and rotational support. At higher resolution, the effect cancels out: a merger that forms a slowly-rotating system in a dry case still forms an equally slow-, or even a bit slower-rotator in the corresponding wet case.

The next section focusses on interpreting the origin of the resolution effect in the formation of slowly-rotating ellipticals. We in particular show that it is not an artifact caused by different initial conditions or a bias in the simulated orbits, but a real effect related to the way the violent relaxation during the merger itself is treated.

4 ORIGIN OF THE RESOLUTION EFFECT

We here show that the above-mentioned discrepancies observed in the simulations that produce slow-rotating ellipticals at high resolu-
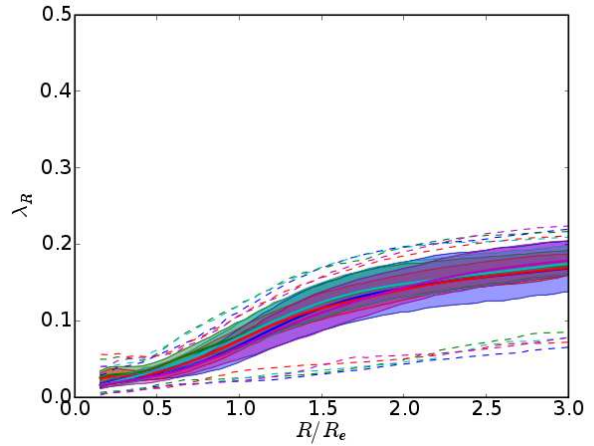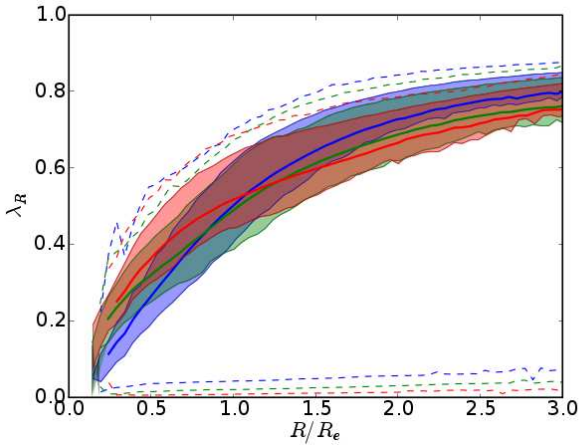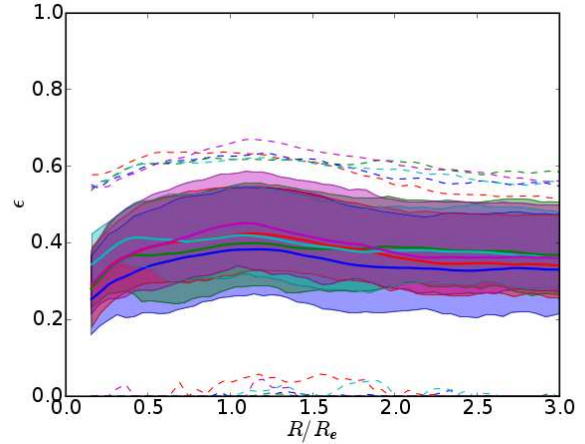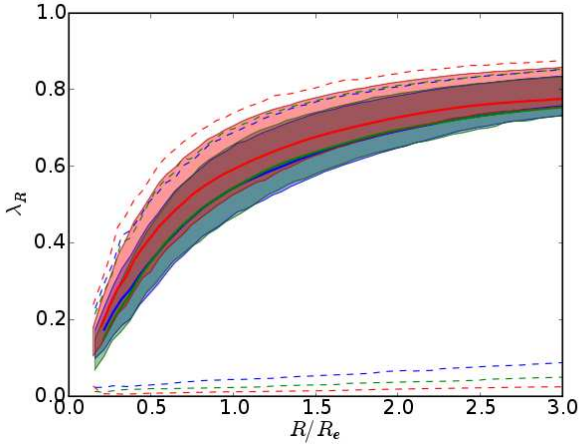
Figure 8. λ_R profiles of the two Wet progenitors as a function of R/R_e. Colours, as in previous Figures, with the *standard-res* in blue, the *high-res* in green, the *very high-res* in red.

Figure 9. ϵ (top) and λ_R (bottom) profiles in function of R/R_e for 5 slightly different orbits for the Wet-Slow-High simulation.

tion, are really attributable to the physical modeling of the merging process. They are not artefacts related to initial conditions of the progenitor galaxies and/or interaction orbits that would vary with the resolution.

4.1 The progenitor galaxies

We first check that the progenitors galaxies are similar at any resolution. To this aim, we analyze their kinematic properties, in particular the λ_R profiles – ϵ and a_4/a parameters are less relevant for disc-dominated galaxies. Since the merger simulations were performed after an isolated relaxation of each progenitor galaxy (see Sect. 2), we analyzed the progenitors from a snapshot right after this relaxation period, so that the results (Figure 8) are representative of the conditions under which the mergers occur.

The two progenitor galaxies have quite similar angular momentum profiles (Figure 8). There are some fluctuations, but they are not systematically corresponding to an increase or decrease of λ_R with resolution. They are also weaker than the discrepancies found in the final merger remnants. Actually, they result for a large part from the effective radius changing slightly with the resolution, and profiles of λ_R as a function of the absolute radius (in kpc) show smaller differences than the profiles in units of the effective radius.

These fluctuations cannot therefore be the main cause for the observed resolution effects in the merger remnants.

4.2 Interaction orbits

Simulations at the three resolutions are started with the same relative position, velocity and inclination for the two interacting progenitors. However, varying the resolution may result in slight differences in dynamical friction and angular momentum exchanges, if these processes are resolved differently, and the interaction orbits might diverge before the merger actually takes place. If this were the case, our results would be attributable to different orbits rather than different treatments of the merging process itself.

We found that the positions at the first pericentric passage vary by 2.1 kpc on average, and the velocities by 9 km s^{-1}. Although these differences seem small and no systematic variation with resolution appeared, we further investigated their potential effect. To this aim we performed four new realizations for the Wet-Slow model at *high-res*, with variations of the position or the velocity twice larger than the average values above (i.e., ± 3.6 kpc and ± 18 km s^{-1}, respectively). The results are shown in Figure 9 for the morphological and kinematic profiles of ϵ and λ_R. Changes are minor and differences arising in the interaction orbits cannot explain the variation of the results with resolution.
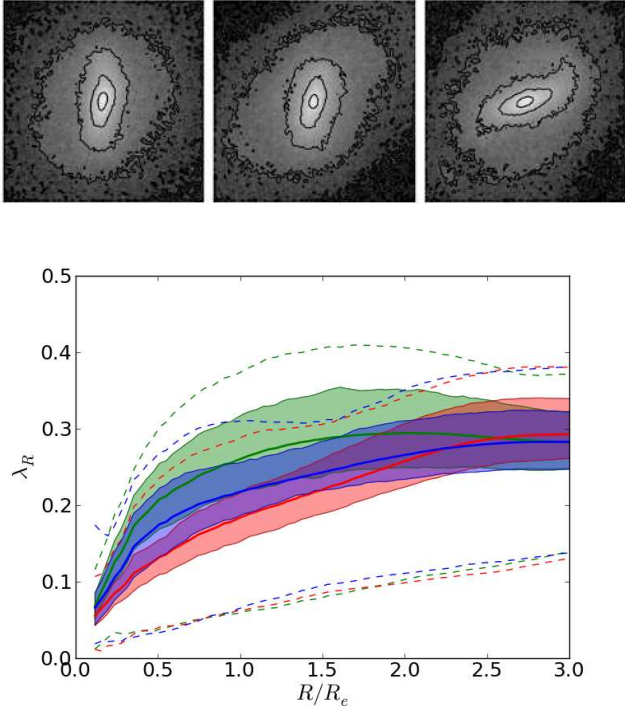
Figure 10. Top panel Stellar (old plus young) intensity maps for three different realisations of the Wet-Slow-Standard simulation. The field of view is 16×16 kpc^2. The simulation used in the study is in the left panel. **Bottom panel** λ_R profiles as a function of R/R_e of the three above simulations.

4.3 Robustness of the Wet-Slow-Standard simulation

As the resolution effect found in the formation of slow rotators, in particular in the Wet-Slow model, cannot be attributed to a change in the initial conditions and interaction orbit, it likely relates to the physical treatment of the merging process itself. Nevertheless we wanted to check whether or not this could still be attributed to the presence of particle noise, which is higher in the *standard-res* cases.

The Wet-Slow model at the *standard-res* shows a strong stellar bar, contrary to the *high-res* and *very high-res* cases. We wanted to make sure that this bar is a robust consequence of the high rotational support of the standard-resolution case, and is not a misinterpreted effect that arised from a particular realization of the particle noise.

To this aim, we performed two other Wet-Slow-Standard simulations with the same initial conditions but different, random realizations of the particle noise. The final stellar distribution, shown in Figure 10 all show a similar bar, and the λ_R profiles are also relatively similar to the original Wet-Slow-Standard model – there are some variations, but the λ_R distributions of the three realizations overlap with each other, and the three systems are equally fast rotators. These two new realisations are also shown in dashed and dotted lines on Figure 2 and 11 and again share common properties with the initial Wet-Slow-Standard model, and hence the same differences compared to the higher-resolution cases.

Thus, the role of bars and spiral patterns in redistributing the mass and angular momentum in the standard-resolution Wet-Slow model is robust, independent of a particular realization of the particle noise. We also find (see next subsection) that the time variations of the gravitational potential during the interaction and merger are similar for the three *standard-res* realizations.
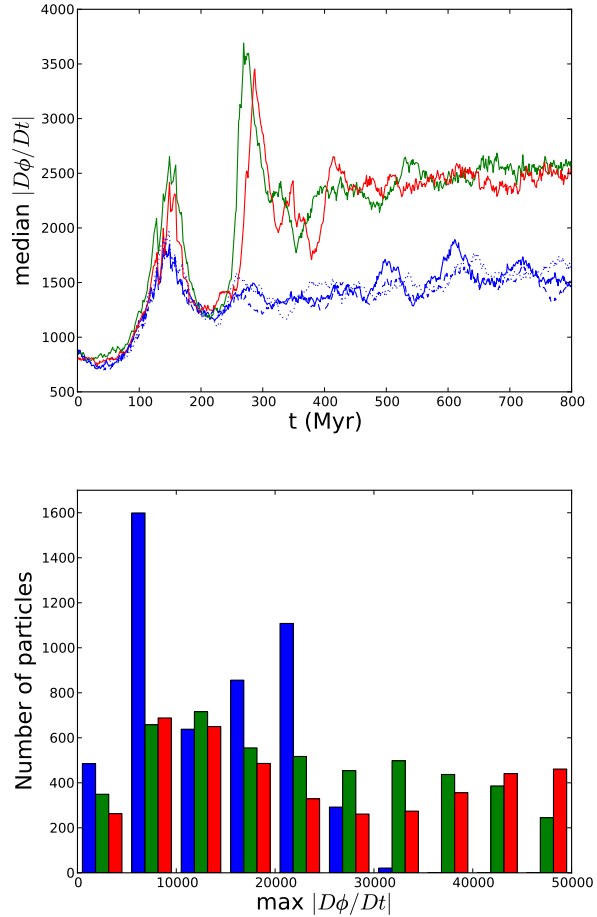
Figure 11. Top panel Median variations of the gravitational potential (in arbitrary units) over 5000 test particles as a function of time. The two other realisations of the Wet-Slow-Standard simulation are shown in dashed line. **Top panel** Histogram of the maximum variation of the gravitational potential (in arbitrary units) of each particle over the simulation.

4.4 The on-going merger

At this point, we have established that the differences observed in the merger remnants do not result from variations in the initial conditions, interaction orbits, or particle noise. The differences should then arise in the physical treatment of the merging process, which would mean that they are "robust" effects, potentially affecting any simulation with any numerical code. Varying the spatial and mass resolution could affect the detailed evolution of the dissipative component (including star forming structures), and this could in turn modify the overall orbital structure of the merger remnant (see Barnes & Hernquist 1996; Naab, Jesseit & Burkert 2006; Cox et al. 2006). However we have seen that the resolution effect does not completely disappear in dry mergers. A more general effect can be the treatment of the violent relaxation, i.e. the rapid changes of gravitational potential that are responsible for the evacuation of energy and angular momentum from the main body of the merger remnant – these quantities being carried away by a low fraction of the mass expulsed at large radii. This process of course plays a more important role in the formation of slow rotators than in the formation of fast rotators. The resolution effects are much more

important for slow rotators than fast ones (Sect. 3), which suggests that they do actually relate to the violent relaxation process.

To quantify the importance of violent relaxation in our merger simulations, we followed, in the Wet-Slow models, the variations of the gravitational potential $D\phi/Dt$ of 5000 randomly chosen "test" particles, all of which are stellar particles existing at $t = 0$, all along the simulations. The derivate is Lagrangian, since it follows the motion of each particle. In an isolated galaxy, $D\phi/Dt$ relates to the variation of potential along the orbit of each particle, in particular their radial excursion in the potential well of the galaxy. During the interaction and mergers, peaks of $D\phi/Dt$ should trace the importance of scattering by local density fluctuations through the violent relaxation process.

The top panel of Figure 11 shows the median value of $|D\phi/Dt|$ as a function of time – we take the absolute value for each particle, as a particle moving inward or outward can be considered with the same behaviour. Before the merging (*i.e* before ∼150 Myr) the three resolutions are identical, meaning that there is no difference in the progenitors during the approach phase, modest values of $|D\phi/Dt|$ simply correspond to modest radial excursions of particle in the progenitor disc galaxies.

After the merger, each simulation shows a relatively constant $|D\phi/Dt|$ in a relaxed system, but the value is higher at *high-res* and *very high-res*, indicating larger radial excursions of stellar particles compared to the *standard-res* case. More radial orbits are indeed expected for slow rotators compared to the *standard-res* fast rotator. We note again that the different orbital structure does not only affect the gas and the young stars formed during the mergers, but also the old stars present before the merger itself (see also Figure 5 and Sect. 3).

During the merging process, a first peak in the median $|D\phi/Dt|$ occurs at the first pericenter passage, after about 150 Myr, but is more pronounced at high(est)-resolution. Another peak is found at the *high-res* and *very high-res* during the final coalescence at $t \sim 280$ Myr , but is much weaker in the *standard-res* case. The final coalescence does take place at the same moment for the three resolutions, but is a smooth process in the *standard-res* case, while it is accompanied by rapid variations of the potential undergone by stellar particles at high resolution. The bottom panel of Figure 11 shows the maximum variation of $|D\phi/Dt|$ for each particle over all the simulation. The distribution at *standard-res* is very different from the distribution at *high-res* and *very high-res*. This confirms that the particles at *standard-res* undergo less rapid variation of the potential, i.e. lower peaks of gravitational forces.

This overall demonstrates that the relaxation process, during the merging of galaxies, is smoother at low resolution than at high resolution. We have shown previously that the *high-res* and *very high-res* simulations resolve much more dense substructures, like gas filaments, stellar clusters, compact cores, etc. Our interpretation is then that these local density peaks are accompanied by rapid variations of the gravitational potential, which scatter the stellar orbits, evacuate the angular momentum, and form, for favorable orbits, slowly-rotating elliptical galaxies. At low resolution, these rapid and local fluctuations of the density and potential are largely missed, hence the merging process is smoother, and more angular momentum remains in the main stellar body of the merger remnant.

Density fluctuations are of course stronger in the dissipative component (gas) and the young stars formed therein, which likely explains why the resolution effect is stronger in wet mergers. Nevertheless, old stars are clearly affected as well, as was shown above.

This also explains why the effect of gas in a wet merger, compared to a dry merger at fixed resolution, is different for

standard-resolution models and high-resolution ones (Sect. 3.2.2). At *standard-res*, the gas remains relatively smooth, promotes the survival/rebuilding of a stellar disc component, thus increasing the rotational support in the final merger remnants. At higher resolution, the presence of gas forms many dense small-scale substructures of gas and young stars (consistent with observations, see e.g. Bournaud, Duc & Emsellem 2008), these substructures increase the degree of relaxation during the merging process, not just for the gas and young stars but also for the old stars. Thus, while the presence of gas should still promote the survival/rebuilding of a disc component in the merger remnant (our high-resolution wet-slow remnant does have a low-mass disc componnent of gas and young star), it also promotes orbital scattering and evacuation of the angular momentum for the whole baryonic mass, but the latter effect is missed if the resolution is too low. This explains why, at high resolution, the merger remnant (in the Wet-Slow case) does not have a higher rotation support or a more prominent disc component than the corresponding Dry-Slow case, an in fact even has a somewhat lower λ_R at one effective radius.

The high-resolution simulations, compared to the standard cases, resolve the formation of dense and relatively massive substructures (clusters, cores, filaments of 10^{5-7} solar masses) that scatter the stellar orbits and evacuate the angular momentum from the main body of the merger remnant. *Very high-res* simulations show a relatively reasonable convergence compared to the *high-res* ones: they resolve the same massive substructures, plus lower mass ones ($\sim 10^{4-5}$ solar masses) that are more numerous but are much less efficient to scatter the orbits and affect the relaxation of the merger remnant, as the corresponding relaxation timescale is much longer. It is thus expected that results converge at a high-enough resolution.

4.5 Timestepping and code specificities

Our results have been obtained with a given code and one can naturally wonder whether or not other codes would show the same resolution effect. We in fact expect no fundamental differences in the output from different codes, given that similar substructures are formed initially: this relaxation effect is mostly gravitational, and this should be treated rather similarly in grid-based and tree-codes. The main question remains then whether or not other codes would form substructures similar to those found in our simulations (e.g., with a similar mass spectrum, Bournaud, Duc & Emsellem 2008): this a priori depends on the modeling of gas cooling and turbulence dissipation processes.

Another specificity of the code employed is its fixed timestep. A small timestep may better resolve the scattering of stellar orbits by dense substructures, in particular at high resolution. This would actually tend to increase the effect of resolution that we have found, which justifies studying the resolution effect at fixed time-step rather than decreasing the timestep at increasing spatial resolution. This way, the effects found can be robustly attributed to the spatial resolution. The time-step itself may have additional, separate effects, in our code or any other, that should be studied separately at fixed (high) resolution.

5 A SYSTEMATIC EFFECT IN THE FORMATION OF SLOW ROTATORS

To ensure that the resolution effect in the formation of slow rotators is a systematic one, and not specific to one simulated orbit, we have
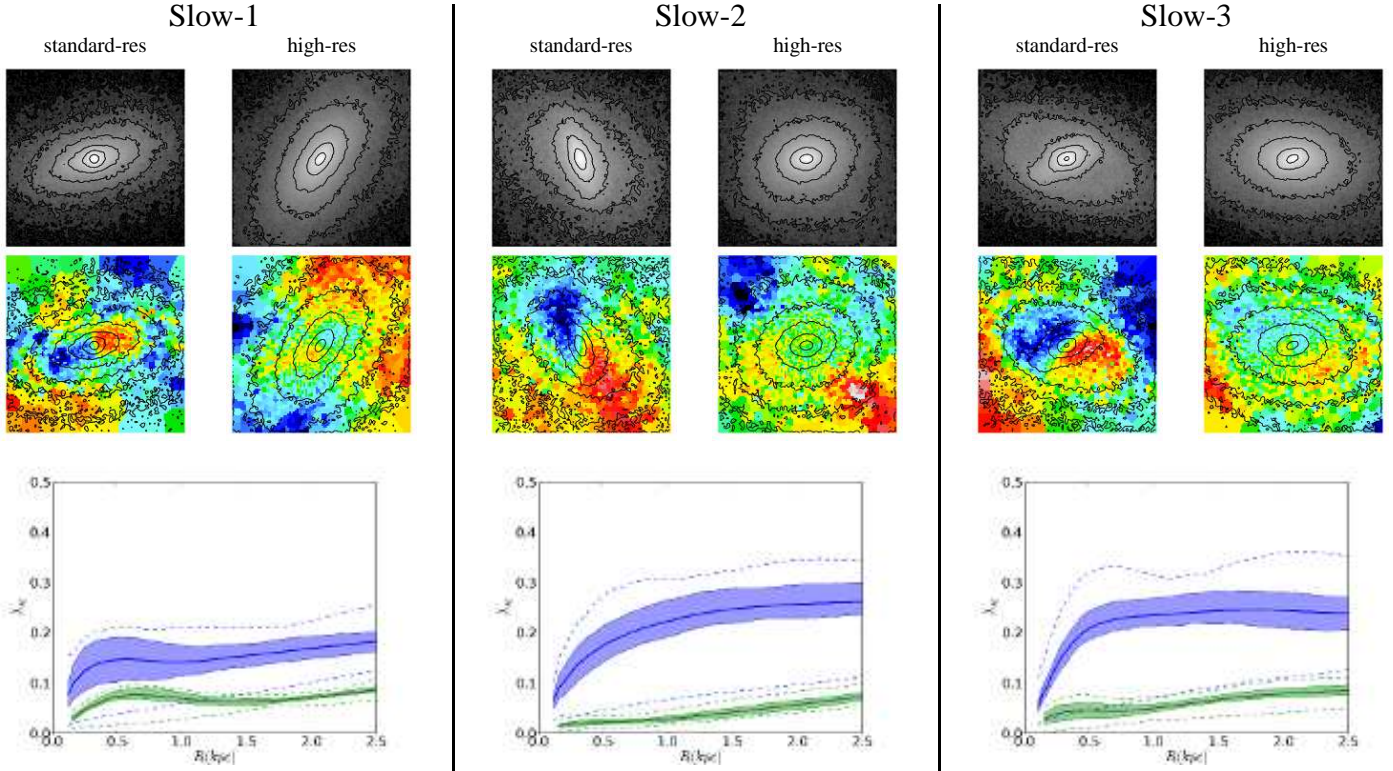
Figure 12. Projected density and velocity fields for the *standard-res* and *high-res* resolutions and their respective λ_R profiles (the *standard-res* in blue, the *high-res* in green) for the *slow-1, slow-2, slow-3* models (from left to right resp.).

selected other major mergers that form slow rotators at high-enough resolution in a larger simulation sample (Bois et al. in preparation), and re-simulated them at lower resolution. These three additional mergers were not simulated at the *very high-res* but at a resolution which is actually a bit higher than the *high-res*, with a spatial resolution of 58 pc and 2×10^6 particles per component par galaxy (i.e. a total of 1.2×10^7 particles). They were also resimulated at the same *standard-res* level as the previous models. The gas fraction in these models is 10%, so as to check that a high gas fraction is not required for the resolution effects to arise. Compared to the original slow-rotator simulation, one orbital parameter is changed in each case. Model *slow-1* has a retrograde orbit for both galaxies, model *slow-2* has a pericenter distance of 25 kpc, model *slow-3* has an orbit inclination of 25 degrees for both discs.

The projected stellar density maps and line-of-sight velocity fields are shown for these models, at standard- and high-resolutions, under the flattest projections in each case, on Figure 12 (the λ_R profiles being also shown in this figure). These three merger remnants are slow rotators (at least at one effective radius) at *high-res*, but at *standard-res* they all have a much higher angular momentum λ_R, and a velocity field aligned with their morphological axis – the high-resolution slow rotators have important kinematic misalignments and KDCs.

These additional cases confirm the effect found and analyzed in detail in the original slow-rotator model, with about the same degree of discrepancy between the standard-resolution and higher-resolution models. Because we have spanned the four slowest rotators among the 1:1 mergers from the Bois et al. (in preparation) sample, the resolution effect seems to strongly affect the modeling of a significant number of slow-rotating ellipticals, if not all, even with modest gas fractions (here 10% of the baryonic mass).

6 DISCUSSION AND CONCLUSION

In this paper, we have studied the effect of numerical resolutions (spatial and mass resolution) on the global properties of merger remnants. Our simulations at "standard" resolution, are comparable to the majority of merger simulation samples published in the last few years: the spatial resolution (gravitational softening and typical hydrodynamical smoothing lengths) is 180 pc, and the number of particles $\sim 10^5$ per galaxy and per component (gas, stars, and dark halo). These simulations have been compared to models of the same mergers with increased resolution, up to 32 pc and almost 10^7 particle per galaxy and per component.

We have analyzed the morphology and kinematics of the relaxed merger remnants. In particular, we have studied whether they are "fast rotators", with an apparent spin parameter $\lambda_R > 0.1$ and have small misalignments between the morphological and kinematic axes, i.e. in broad terms early-type galaxies with significant flattening and rotational support. At the opposite end, "slow rotators" are systems with a low $\lambda_R \leqslant 0.1$ (at one effective radius), large kinematic misalignments, i.e. early-type galaxies dominated by (anisotropic) pressure support and low residual rotation. Such slow rotators usually have central KDCs in our high-enough resolution simulations.

Our mains findings can be summarized as follows:

• The formation of *fast-rotators* is not significantly affected by numerical resolution. Models that produce fast rotators at the highest resolution also result in fast rotators at lower resolution, with some random fluctuation of their properties but no sign of systematic variation in the morphology or angular momentum profile against resolution.

• The formation of *slow-rotators* is greatly affected by numeri-

cal resolution. Models that produce slow rotators at the highest resolution result in much faster rotators at lower, standard resolution. The effect is present, but relatively minor, in purely collisionless dry mergers. Discrepancies become major in wet mergers, even in cases with modest gas fractions like 10% of the baryonic mass.

• These effects cannot be attributed to our choice of initial conditions or interaction orbits, but actually relate to the physical treatment of the merging process itself. In particular, small-scale density fluctuations increase at high resolution, and they participate to scattering stellar orbits and largely influence the final degree of relaxation and orbital structure in the merger remnants.

• The effect of gas on the properties of merger remnants is generally considered to consist in preserving a higher angular momentum, in particular through enhancing the survival/rebuilding of disc components in merger remnants. We find that this picture is incomplete: at high resolution, gas still reforms discy components, but also forms a large number of dense substructures (massive star clusters, dense nuclei, tails and filaments, etc) that trigger rapid variations of the gravitational potential and the degree of relaxation of the final system. This effect is missed with a too low resolution. At high-enough resolution, adding gas to a given merger does not necessarily increase the rotational support of the final merger remnant; we even find a case of a wet merger with 20% of gas that has a final angular momentum parameter λ_R slightly lower than the corresponding dry merger.

At the present stage, our results do not indicate how frequently real slowly-rotating ellipticals were formed by binary (wet) mergers of disc galaxies, but they show that this can be a robust pathway for their formation. In the course of the ATLAS^{3D} project, we are conducting, analyzing, and comparing a large set of numerical simulations for various formation mechanisms, in order to derive which is (are) the main formation mode(s) of real slow rotators in the nearby Universe. Our present results already indicate the limitations of existing samples of galaxy merger simulations, and will then serve to estimate the required resolution, the limitations of numerical models and their possible biases.

More generally, the immediate implications of these findings on our understanding of early-type galaxy formation are:

• High resolution in simulations of major mergers does not just allow to resolve small-scale structures like nuclear systems and star clusters, but impacts the whole global properties of the elliptical-like merger remnants, at least for the slow-rotating ones.

• The formation of slow-rotating elliptical galaxies can be achieved through a major merger relatively more easily than previously believed. It can be frequent even in wet mergers with relatively high gas fraction, and with late-type, disc-dominated progenitor galaxies.

• Repeated mergers and/or dry mergers of galaxies that are already early-type systems are thus not the only theoretical path to produce slow-rotating galaxies. Major mergers of two dic galaxies, including wet mergers, can produce slow-rotating early-type galaxies. Further studies are needed to determine how common this formation mechanism is for slow-rotators.

• Quantitative comparisons of major merger simulation results with the observed properties of real early-type galaxies require high resolution models. A typical requirement, according to our study, would be a spatial resolution better than 100 pc for both the gravitational N-body aspects (i.e., softening length) and the hydrodynamical ones (for instance, the size of groups of particles other with quantities are smoothed in SPH models). The mass resolution should correspond to at least $\sim 10^6$ particle per galaxy per compo-

nent, which typically corresponds to a mass resolution $\sim 10^4$ M$_\odot$ for the gas discs of bright spiral galaxies. We find reasonable convergence above this resolution, but cannot rule out that some systematic effects still exist; in any case simulations below this resolution level show clear and strong resolution effects. Unfortunately, most published samples of major merger simulations are below this typical resolution limit.

• The small-scale properties of interstellar gas and clustered star formation are important for the global, large-scale properties of merger remnants. Simulations directly resolving gas cooling down to low temperatures, the formation of cold (molecular) gas clouds and star formation therein, are highly desirable to understand the whole process of early-type galaxy formation. Modern hydrodynamic codes are promising in this respect (e.g. Bournaud, Elmegreen & Martig 2009; Kim, Wise & Abel 2009).

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APPENDIX A: ADDITIONAL MORPHOLOGICAL AND KINEMATIC PARAMETERS FOR THE MERGER REMNANTS

In this appendix, we provide further details on the morphology and kinematics of the four merger remnants analysed in the present paper.

A1   Morphology

The ellipticity ϵ and a_4/a profiles are shown in Figure A1 as a function of R/R_e. The apparent differences sketched in Sect. 3.1 are confirmed quantitatively in the radial ellipticity profiles. Within 1.5 R_e, there are small differences in the Dry/Wet-Fast and Dry-Slow simulations. The ellipticity outside 1.5 R_e is however quite similar at all resolutions for these three simulations: the minimum ellipticity is basically 0, the mean is 0.33 ± 0.03 and the maximum is 0.55 ± 0.05 at 2 R_e for all three *standard-res*, *high-res*, *very high-res*.

The Wet-Slow simulation shows much larger differences. Outside 0.5 R_e, the *high-res* and *very high-res* are similar. The ellipticity profile of the *standard-res* has then a completely different appearance: between 0.6 and 2 R_e, 75% of the projections have an ellipticity higher than 0.4, and the reached maximum in ϵ is 0.75 (versus $\sim$ 0.6 for the other two resolutions). In the outer part ($R > 2R_e$) the ellipticity of most of the projections is decreasing but its maximum is still larger than 0.7.

The same trends are observed in the a_4/a profiles. In the Dry/Wet-Fast and Dry-Slow simulations at all three resolutions, the mean a_4/a is around 0. Then 50% of the projections are between -2 and 2%. The *high-res* Wet-Fast and the *very high-res* Dry-Fast simulations are only slightly more boxy. The a_4/a profile of the Wet-Slow simulation dramatically confirm what we observe for the ellipticity. The *standard-res* clearly departs from the *high-res* and *very high-res*, which are quite similar. Between 0.5 and 1 R_e, the

projections of the *standard-res* span a very large range of a_4/a. Between 1 and 2 R_e, 75% of the projections have a discy shape, and the isophotes of the merger remnant become then increasingly boxy going outwards.

A2   Kinematics

The Wet-Slow simulation has been treated in the paper, we will thus focus on the three other simulations.

Left panels of Figure A2 show the velocity and velocity dispersion curves for the mean-ellipticity projection along the global kinematic position angle. In the Dry-Fast simulation, the central slope of the rotation curve at *very high-res* is slightly shallower, and the dispersion about 15% smaller, but in the outer part the velocity amplitude is similar at all three resolutions, with a velocity of about 60 km s^{-1} at 6 kpc, and dispersions values going to about 150 km s^{-1}. The Wet-Fast simulations show consistent velocity profiles at all three resolutions, with a velocity amplitude of 60 km s^{-1} at 6 kpc, and dispersion decreasing outwards down to $\sim$ 125 km s^{-1}. Again, the velocity curves for the Dry-Slow simulation are all very similar, but these profiles clearly reveal the previously observed KDC which appears here as a kpc-size core counter-rotating with respect to the outer part. Note the *standard-res* dispersion profile which is about 10% smaller than for the other two higher resolutions.

In right panels of Figure A2, we now compare the simulations using the apparent angular momentum λ_R. These figures clearly show that the Dry/Wet-Fast simulations (top and second from top) both result in *fast-rotators*, the mean values of λ_R is 0.2 and the maximum about 0.25 at 1 R_e for the three resolutions. This confirms our previously mentioned results that the spatial and mass resolutions do not seem to have a significant effect on these merger remnants.

The analysis of the morphology and kinematics of the Dry-Slow simulation did show mild differences in the remnants for varying resolutions, the λ_R profiles exacerbate these small discrepancies. At *standard-res*, λ_R is an increasing function of radius, with 75% of all projections having values below 0.1 at 1 R_e, and 25% above 0.1. However, if we are taking into account the projection which maximises λ_R, the *standard-res* remnant should be classified as a *fast-rotator*. In the same context, both the *high-res* and *very high-res* are classified as *slow-rotators*. They have not the same profiles but have a similar overall behaviour: λ_R first increases up to about 0.5 R_e, and then decreases (up to 1.5 R_e for the *high-res* and 1 R_e for the *very high-res*). Outside 1.5 R_e, λ_R increases again outwards. Such a λ_R profile is the clear signature of large-scale KDCs as mentioned in Emsellem et al. 2007 (see also McDermid et al. 2006).
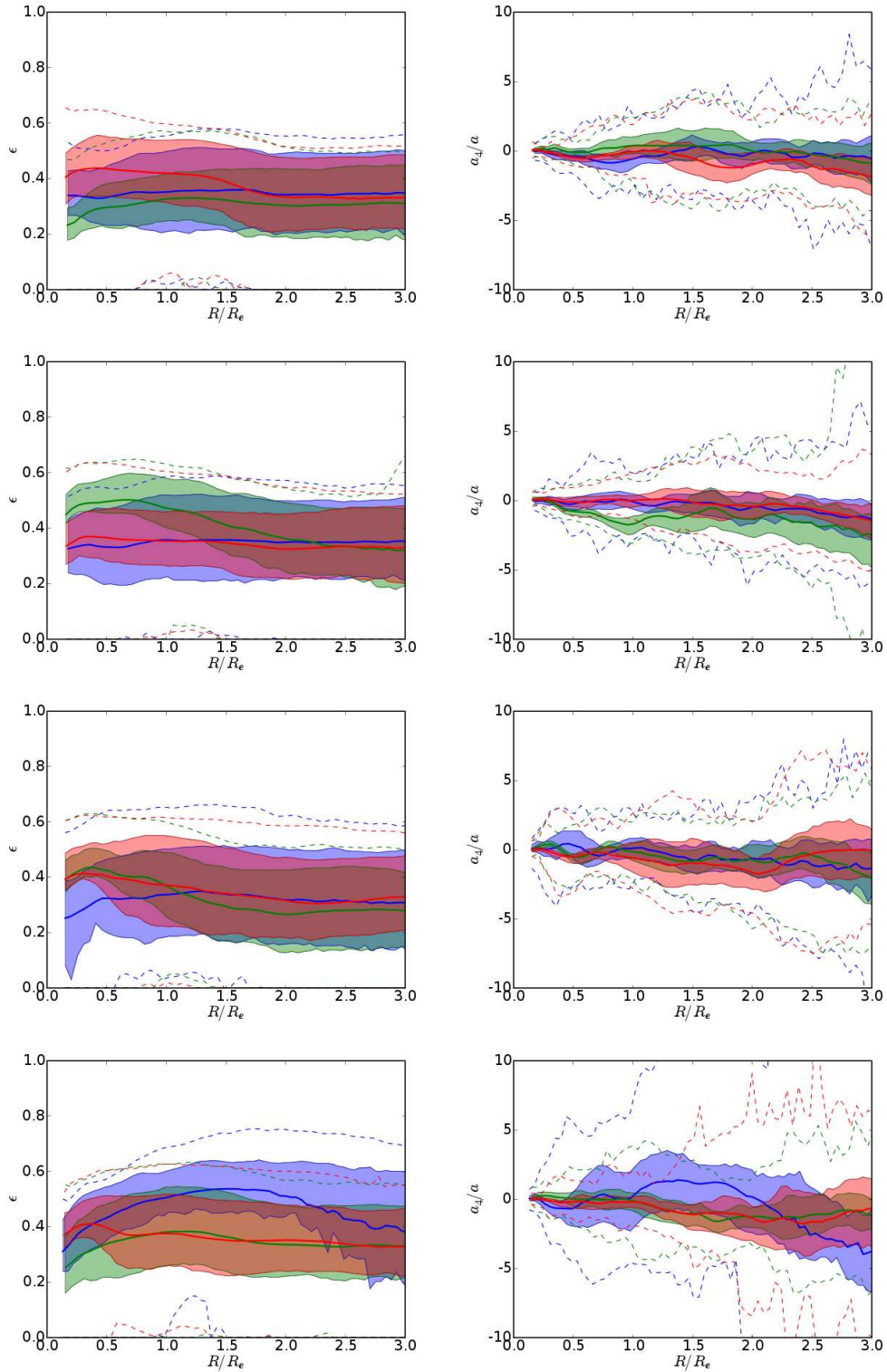
Figure A1. Ellipticity and a_4/a profiles (left and right panels resp.) of the four simulations as a function of R/R_e. From top to bottom : Simulations Dry-Fast, Wet-Fast, Dry-Slow and Wet-Slow. The three resolutions are shown with different colours : the *standard-res* in blue, the *high-res* in green and the *very high-res* in red. For each resolution, we plot five lines which correspond to the minimum and maximum at each radii (dashed lines), the mean value (thick solid lines) and the first and third quartiles (thin solid lines). The inter-quartiles space (which correspond to 50% of all projections) is filled with the colour associated with the resolution.
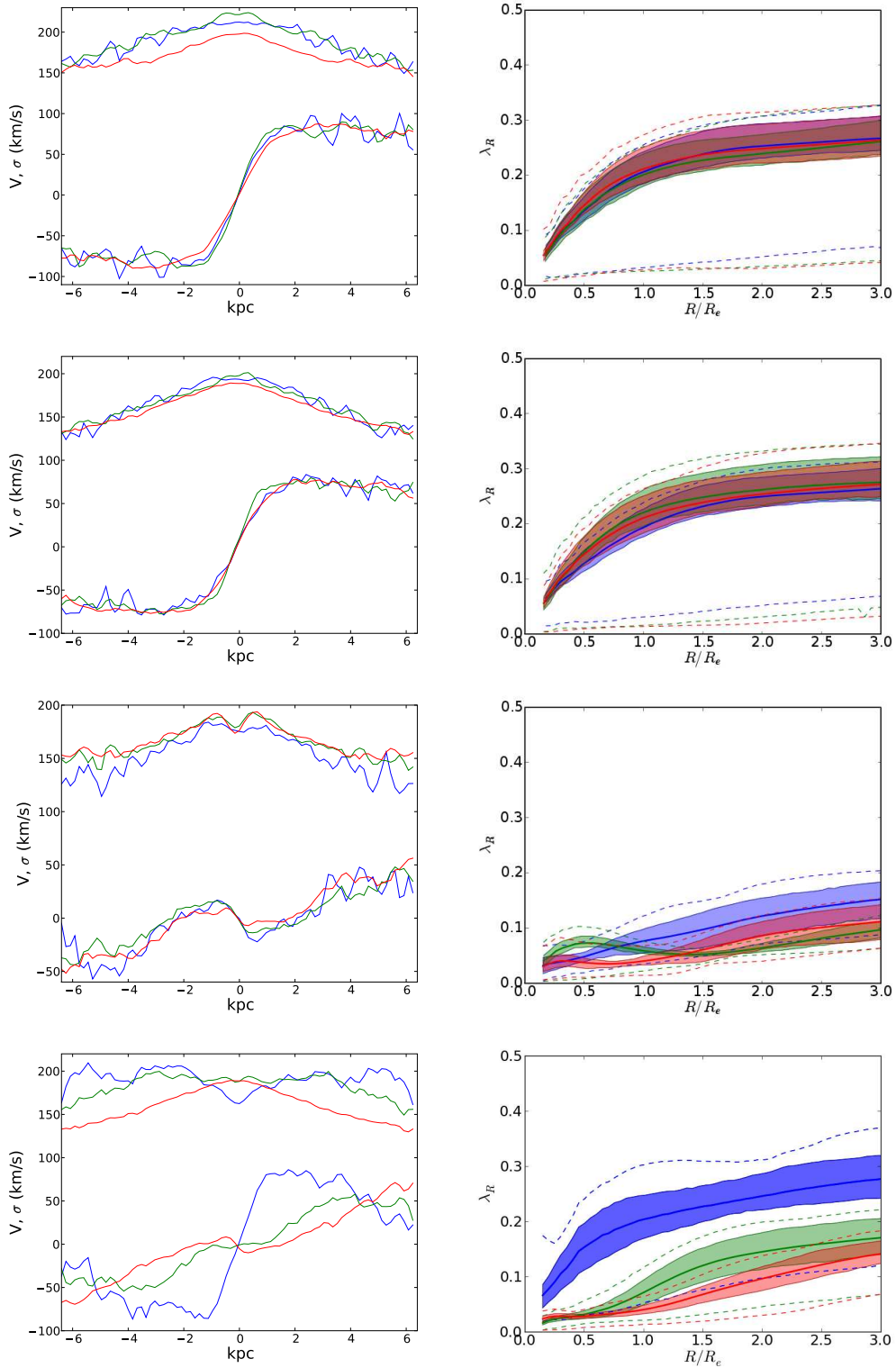
Figure A2. Left Panels The radial velocity and velocity dispersion profiles (in km s^{-1}) for the mean ellipticity projection along the global kinematic position angle (radius in kpc). **Right Panels** λ_R profiles as a function of R/R_e. From top to bottom : Simulations Dry-Fast, Wet-Fast, Dry-Slow and Wet-Slow. The three resolutions are shown with different colours : the *standard-res* in blue, the *high-res* in green and the *very high-res* in red.