WISDOM project – V. Resolving molecular gas in Keplerian rotation around the supermassive black hole in NGC 0383

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ABSTRACT

As part of the mm-Wave Interferometric Survey of Dark Object Masses (WISDOM), we present a measurement of the mass of the supermassive black hole (SMBH) in the nearby early-type galaxy NGC 0383 (radio source 3C 031). This measurement is based on Atacama Large Millimeter/sub-milimeter Array (ALMA) cycle 4 and 5 observations of the ^{12}CO(2–1) emission line with a spatial resolution of $58 \times 32\,\mathrm{pc}^2$ ($0\rlap.{''}18 \times 0\rlap.{''}1$). This resolution, combined with a channel width of $10\,\mathrm{km\,s^{-1}}$, allows us to well resolve the radius of the black hole sphere of influence (measured as $R_{\mathrm{SOI}} = 316\,\mathrm{pc} = 0\rlap.{''}98$), where we detect a clear Keplerian increase of the rotation velocities. NGC 0383 has a kinematically-relaxed, smooth nuclear molecular gas disc with weak ring/spiral features. We forward-model the ALMA data cube with the Kinematic Molecular Simulation (KinMS) tool and a Bayesian Markov Chain Monte Carlo method to measure a SMBH mass of $(4.2 \pm 0.7) \times 10^9\,\mathrm{M_\odot}$, a F160W-band stellar mass-to-light ratio that varies from $2.8 \pm 0.6\,\mathrm{M_\odot/L_{\odot,F160W}}$ in the centre to $2.4 \pm 0.3\,\mathrm{M_\odot/L_{\odot,F160W}}$ at the outer edge of the disc and a molecular gas velocity dispersion of $8.3 \pm 2.1\,\mathrm{km\,s^{-1}}$(all 3σ uncertainties). We also detect unresolved continuum emission across the full bandwidth, consistent with synchrotron emission from an active galactic nucleus. This work demonstrates that low-J CO emission can resolve gas very close to the SMBH ($\approx 140\,000$ Schwarzschild radii) and hence that the molecular gas method is highly complimentary to megamaser observations as it can probe the same emitting material.

Key words: galaxies: Individual NGC0383 – galaxies: kinematics and dynamics – galaxies: nuclei – galaxies: ISM – galaxies: Early type

1 INTRODUCTION

Early-type galaxies, as gravitationally-bound stellar systems, lie on a tight "Fundamental Plane" defined by their mass (luminosity), size (half-light radius) and second ve-

locity moment (velocity dispersion; e.g. Djorgovski & Davis 1987; Dressler et al. 1987). Late-type galaxies follow less tight correlations such as the Tully-Fisher relation between mass (luminosity) and rotation velocity (Tully & Fisher 1977; see Courteau et al. 2014 and Section 4 of Cappellari 2016 for reviews of the fundamental planes of galaxies). Comparing central supermassive black hole (SMBH) mass

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measurements with these galaxy properties has revealed further relations connecting, for example, bulge mass, stellar mass (or luminosity) or Sersic concentration index to the SMBH mass (Kormendy & Richstone 1995; Magorrian et al. 1998; Marconi & Hunt 2003; Haering & Rix 2004; Graham 2007). This has led to the prevailing theory that SMBHs, despite their comparatively small masses, are a major influence on galaxy evolution (e.g. Beifiori et al. 2012; van den Bosch 2016). However, these relations are poorly constrained, with relatively few data points drawn from biased samples, and with large uncertainties. Further evidence indicates both SMBH mass growth at the same rate as (e.g. Mullaney et al. 2012; Madau & Dickinson 2014) and SMBH feedback quenching of (Bundy et al. 2008) star formation by SMBHs. Furthermore, whether all galaxies follow the same relations or not is still inadequately tested. In particular, there is evidence that low-mass late-type and high-mass early-type galaxies follow different co-evolutionary relationships (e.g. McConnell & Ma 2013). Kormendy & Ho (2013) give a comprehensive review of the current state of these relations.

One of the tightest relationships is that between the SMBH mass (M_{BH}) and the stellar velocity dispersion (σ_*; e.g. Gebhardt et al. 2000; Ferrarese & Merritt 2000), but there is again growing evidence of divergence between galaxies of different morphological types or masses (e.g. McConnell & Ma 2013; van den Bosch 2016, esp. their Fig. 2; Krajnović et al. 2018). To fully analyse the extent of the co-evolution between all these galaxy properties, it is essential to gather a larger, more diverse, sample of *reliable* SMBH mass estimates (van den Bosch 2016).

Reliability is achieved by directly measuring the SMBH masses through their gravitational influence. Methods to measure SMBH masses dynamically include observing and modelling the stellar kinematics (e.g. Dressler & Richstone 1988; Cappellari et al. 2002; Krajnović et al. 2009), ionised gas kinematics (e.g. Ferrarese et al. 1996; Sarzi et al. 2001; Walsh et al. 2013) and megamaser kinematics (e.g. Herrnstein et al. 1999; Miyoshi et al. 1995; Greene et al. 2010). However, each of these methods can only be used in a small fraction of the galaxy population, as each is biased towards particular morphologies. For instance, stellar kinematics are often hampered by dust contamination and require either resolving individual stars directly or strong absorption lines in integrated spectra. Megamasers probe material very close to the SMBHs but require an edge-on view and are very rare (being present in only $\approx 5\%$ of objects searched; Lo 2005). They are typically found in Seyfert 2 and low-ionisation nuclear emission region (LINER)-type nuclei of low-mass galaxies. Overall the current sample is biased towards nearby, high surface brightness objects. A new method of measuring SMBH masses is thus required to diversify the sample.

To expand the current sample to all morphological types, galaxy masses and both active and non-active galaxies, our mm-Wave Interferometric Survey of Dark Object Masses (WISDOM) is using a new method exploiting molecular gas observations to trace the velocity fields surrounding SMBHs. The first use of this method with Carbon Monoxide (CO) was by Davis et al. (2013b). SMBH mass measurements in fast-rotator early-type galaxies (Onishi et al. 2017; Davis et al. 2017, 2018), galaxies with ir-

regular gas distributions (Smith et al. 2019), and in the first late-type galaxy with the dense molecular gas tracers HCN and HCO^+ (Onishi et al. 2015) have been successful. Barth et al. (2016a,b) and Boizelle et al. (2019) also used CO to constrain the SMBH mass in the early-type galaxy NGC 1332 and NGC 3258. Most recently, Combes et al. (2019) used CO(3–2) observations to investigate the molecular tori around 7 SMBHs and therefore measure their SMBH masses. Nagai et al. (2019) observed the radial filaments of NGC 1275 in CO(2–1), detecting a rotating disc allowing them to make a SMBH mass estimate, that agrees with the estimate from H_2 observations by Scharwächter et al. (2013).

All these observations can detect the dynamical influence of the SMBH if, as shown in Davis (2014), they have a minimum spatial resolution of approximately two times the radius of the sphere of influence (R_{SOI}) of the SMBH,

$$R_{SOI} \equiv \frac{GM_{BH}}{\sigma_*^2}, \tag{1}$$

where G is the gravitational constant. The use of molecular gas, specifically ^{12}CO, reduces the selection biases normally associated with dynamical SMBH mass measurements, because of the wide range of objects with suitable molecular gas discs, and because the high angular resolution required is easily reached by modern interferometers, e.g. the Atacama Large Millimeter/sub-millimeter Array (ALMA). Indeed, molecular discs are found around the centres of galaxies of all morphological types (e.g. Regan et al. 2001; Alatalo et al. 2013). Furthermore, with rotational transitions in the millimetre/sub-millimetre wavebands, CO is observable without dust attenuation.

NGC 0383 (radio source 3C 031; Edge et al. 1959; Bennett 1962) is a well-known radio galaxy (implying the presence of a large SMBH), it has a very regular central dust disc and it is also strongly detected in CO with a clear double-horned profile (Lim et al. 2000; Okuda et al. 2005; Ocaña Flaquer et al. 2010). This work presents a measurement of the SMBH mass in this galaxy using ALMA observations of the ^{12}CO(2–1) line with a spatial resolution of $58 \times 32 \, pc^2$ ($0\rlap{.}''18 \times 0\rlap{.}''1$). In Section 2, we present the target, observations and data reduction. In Section 3, we describe the dynamical modelling and SMBH mass measurement method. A discussion of the uncertainties and comparisons to other SMBH mass measurement techniques are presented in Section 4. We conclude briefly in Section 5.

2   TARGET: NGC 0383

NGC 0383 is a dusty lenticular galaxy at a distance of 66.6 ± 9.9 Mpc (Freedman et al. 2001). It is the brightest galaxy of its group (the NGC 0383 group), part of the Pisces-Perseus Supercluster (Hudson et al. 2001). NGC 0383 hosts a radio-loud active galactic nucleus (AGN) with spectacular radio jets. The coincident radio source is catalogued as 3C 031 (Edge et al. 1959; Bennett 1962). Observations of the jets are presented in Macdonald et al. (1968), Bridle & Perley (1984), Laing & Bridle (2002) and van Velzen et al. (2012), while the flat-spectrum radio source is described in Healey et al. (2007).

We can estimate the required angular resolution by

predicting the $R_{\rm SOI}$ using the SMBH mass upper limit of Beifiori et al. (2009) ($M_{\rm BH} = 1.1 \times 10^9\,{\rm M}_\odot$, corrected to our assumed distance and inclination, see Section 4.4) and σ_* as listed in van den Bosch (2016) ($\sigma_{\rm e} = 239 \pm 16\,{\rm km\,s^{-1}}$ i.e. σ_* within $R_{\rm e}$ from Beifiori et al. 2009, corrected following Jorgensen et al. 1995). Using Eq. 1 with these values, we obtain $R_{\rm SOI} = 82 \pm 15\,{\rm pc}$, indicating we need an angular resolution of better than $\approx 0\rlap{.}''5$ to attempt to detect the dynamical influence of the SMBH (i.e. to resolve $2R_{\rm SOI}$; Davis 2014).

There are existing $^{12}{\rm CO}(1{-}0)$ observations of NGC 0383 from Lim et al. (2000), Okuda et al. (2005) and in particular single dish observations from the Thorough ANalysis of radio-Galaxies Observation project (TANGO; Ocaña Flaquer et al. 2010). They report the total molecular gas mass enclosed by the Institut de Radioastronomie Millimétrique (IRAM) 30-m telescope beam to be $M_{\rm H_2} = (1.7 \pm 0.2) \times 10^9\,{\rm M}_\odot$. We correct this from 71.06 Mpc to our assumed distance of 66.6 Mpc, yielding $(1.49 \pm 0.19) \times 10^9\,{\rm M}_\odot$ as the total molecular gas mass.

2.1 ALMA Observations

The $^{12}{\rm CO}(2{-}1)$ line in NGC 0383 was observed with ALMA on the 21st of June 2016 at moderate resolution ($0\rlap{.}''5$) and then on the 16th of August 2017 at high resolution ($0\rlap{.}''1$), both as part of the WISDOM project (programmes 2015.1.00419.S and 2016.1.00437.S). Configurations C36-5 (baselines 15–704 m) and C40-8 (baselines 21–3637 m) were used to achieve sensitivity to emission on scales up to $4''$, with on-source integration times of 2.22 and 28.8 min, respectively. A 1850 MHz correlator window was placed over the CO(2–1) line and centred at 226.6 GHz, yielding a continuous velocity coverage of $\approx 2000\,{\rm km\,s^{-1}}$ with a raw channel width of $\approx 1.3\,{\rm km\,s^{-1}}$, fully covering and well resolving the line. Three additional low spectral resolution correlator windows were included to detect continuum emission, each of 2 GHz width.

The raw ALMA data were calibrated using the standard ALMA pipeline, as provided by the ALMA regional centre staff. The amplitude and bandpass calibrator used in the two observations was respectively J0237+2848 and J2253+1608. The phase calibration used J0057+3021 and J0112+3208, respectively, to determine and therefore correct atmospheric phase offsets.

We then used the Common Astronomy Software Applications (CASA; McMullin et al. 2007) package to combine the two configurations and image the resultant visibilities. A three-dimensional RA-Dec-velocity data cube was produced with a binned channel width of $10\,{\rm km\,s^{-1}}$. To balance spatial sampling and resolution, pixels of $0\rlap{.}''035 \times 0\rlap{.}''035$ were chosen, yielding approximately 5 pixels across the synthesised beam major axis.

The data presented here were produced using Briggs weighting with a robust parameter of 0.5, yielding a synthesised beam full-width-at-half-maximum (FWHM) of $\theta_{\rm maj} \times \theta_{\rm min} \approx 0\rlap{.}''18 \times 0\rlap{.}''1$ at a position angle of $6\rlap{.}°6$. The corresponding spatial resolution is $\approx 58 \times 32\,{\rm pc^2}$, so that the predicted $R_{\rm SOI}$ is well resolved with about 2 synthesised beams, i.e. $R_{\rm SOI}/\sqrt{\theta_{\rm maj} \times \theta_{\rm min}} = 1.9$. Continuum emission was detected, measured over the full line-free bandwidth, and then subtracted from the data in the uv–plane using the CASA task uvcontsub. The achieved continuum root mean square

(RMS) noise is $35\,\mu{\rm Jy\,beam^{-1}}$. The continuum-subtracted dirty cube was cleaned (see Högbom 1969 for the CLEAN procedure) in interactively-identified regions of source emission in each channel, to a threshold equal to the RMS noise of the dirty channels. The clean components were then added back and re-convolved using a Gaussian beam of FWHM equal to that of the dirty beam. This produced the final, reduced and fully calibrated $^{12}{\rm CO}(2{-}1)$ data cube of NGC0383, with a RMS noise level of $0.4\,{\rm mJy\,beam^{-1}}$ in each $10\,{\rm km\,s^{-1}}$ channel.

2.2 Line Emission

The final data products used in this paper were created from the clean, fully calibrated data cube. Zeroth moment (integrated intensity), first moment (mean velocity) and second moment (velocity dispersion) maps were created using a masked moment technique (e.g. Dame 2011). The mask was generated by taking a copy of the clean cube and smoothing it, first spatially using a Gaussian with FWHM equal to that of the synthesised beam, and then Hanning-smoothing in velocity. The mask selects pixels with an amplitude in the smoothed cube greater than 0.8 times the RMS of the unsmoothed data cube. The moments, shown in Figs. 2 and 3, are made from the original un-smoothed cube with the mask applied. We note that the masking procedure is only used when creating the moment maps, whilst the fitting is performed on the whole unmasked cube.

A regularly rotating and symmetric molecular gas disc is clearly detected, with no evidence that the disc is disturbed by the strong AGN jets. It extends $\approx 4'' \times 6''$ in projection ($\approx 1400 \times 1600\,{\rm pc^2}$). There is a slight dip in flux at the centre of the zeroth moment, partially due to our masking procedure removing low surface brightness emission spread over a large number of channels close to the central SMBH. This hole becomes much less significant when a simple clipping procedure is used, although this does increase the noise. The enhanced velocities around the centrally-located SMBH are obvious in both the first moment map and the major-axis position velocity diagram (PVD; Fig. 4), the latter constructed by summing pixels within a 5-pixel wide ($0\rlap{.}''175$) pseudo-slit at a position angle of $142°$. The position angle used here and derived from the CO observations is reasonably consistent with the optical position angle as listed in the NASA/IPAC Extragalactic Database (NED)[1] of $\approx 150°$. The moment one colourbar and PVD right-hand velocity axis are the observed line-of-slight velocity minus the systemic velocity of $V = 4925 \pm 4\,{\rm km\,s^{-1}}$ derived in Section 3.2. The large velocity dispersion observed at the centre of Fig. 3 is primarily due to beam smearing.

Fig. 5 shows the $^{12}{\rm CO}(2{-}1)$ integrated spectrum, made by integrating over a $6'' \times 6''$ area of the clean cube, thus encompassing the entire disc. It clearly shows the double-horn shape of a rotating disc, as also observed by Lim et al. (2000) in both CO(1–0) and CO(2–1) and Okuda et al. (2005) in CO(1–0) only. The total CO(2–1) flux is $87.1\,{\rm Jy\,km\,s^{-1}}$.

For this work we want high spatial resolution data and hence use long-baseline interferometric observations. Incomplete uv–plane coverage can thus lead to some flux being
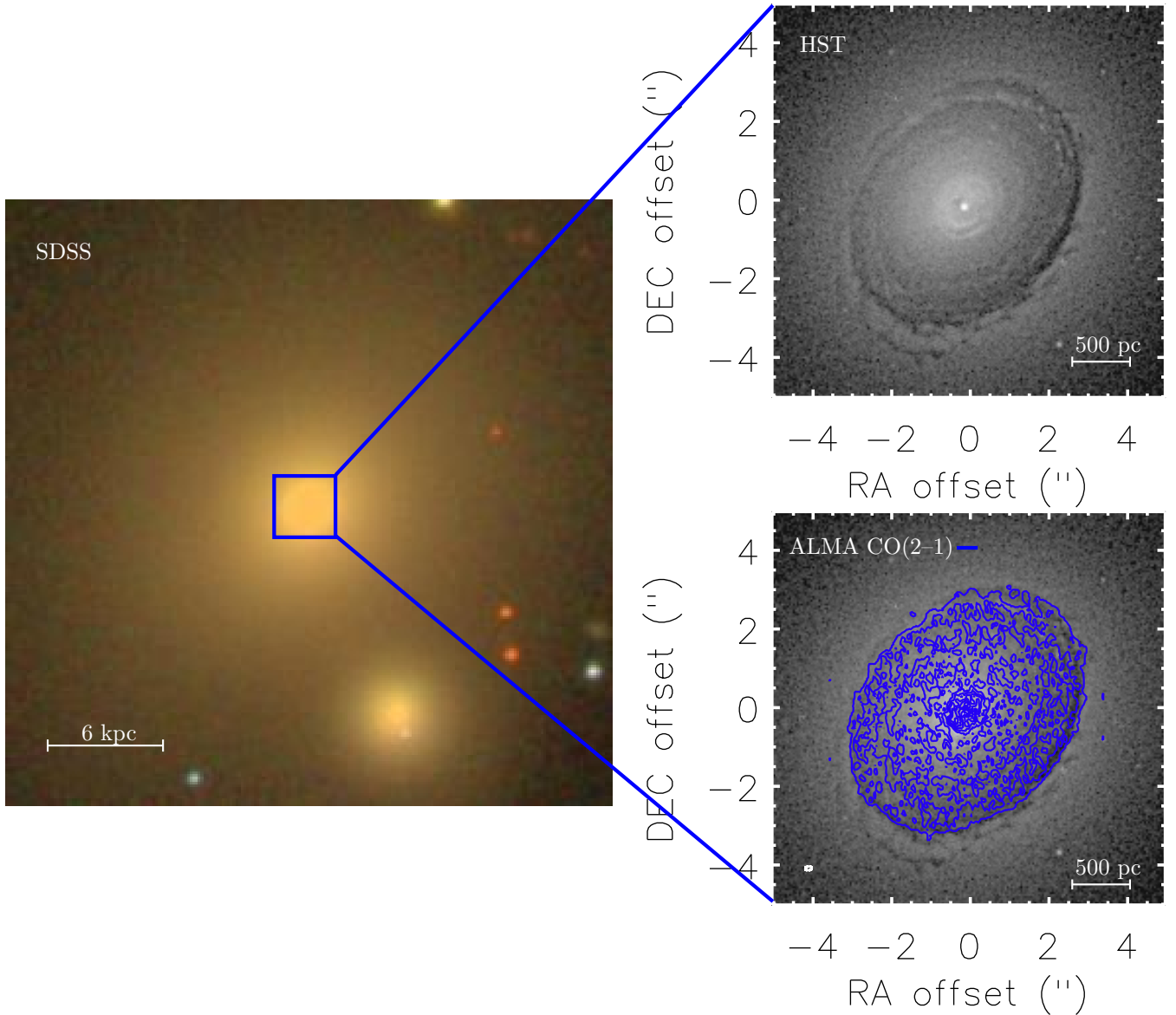
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[1] https://ned.ipac.caltech.edu

Figure 1. *Left panel:* SDSS three-colour (*gri*) image of NGC 0383, $90'' \times 90''$ (29 kpc $\times$ 29 kpc) in size. *Right panel, top:* Unsharp-masked Hubble Space telescope (*HST*) Wide-Field Planetary Camera 2 (WFPC2) F555W image of a 3.2 kpc $\times$ 3.2 kpc region around the nucleus (indicated in blue in the left panel), revealing a clear central dust disc. *Right panel, bottom:* As above, but overlaid with blue ^{12}CO(2–1) integrated intensity contours from our ALMA observations. The synthesised beam ($0.''18 \times 0.''1$ or 58×32 pc^2) is shown as a (very small) white ellipse in the bottom-left corner. The molecular gas disc coincides with the dust disc.

resolved out. To check the scale of this problem, we compare the integrated flux derived from our CO(2–1) observations with that of Ocaña Flaquer et al. (2010), obtained with the 30-m IRAM single-dish telescope. Their CO(2–1) flux is 74.4 ± 2.8 Jy km s^{-1}. As we retrieve slightly more flux than this, and the entire molecular gas disc of NGC 0383 fits within the primary beam of the 30-m telescope, it is unlikely that we resolve out flux in our observations. The lower flux of the single-dish observations may be due to pointing and/or flux calibration errors.

Comparing our CO(2–1) flux of 87.1 Jy km s^{-1} to that of the CO(1–0) line (29.8 Jy km s^{-1}; Ocaña Flaquer et al. 2010), we find a CO(2–1)/CO(1–0) ratio of 0.73 after converting

to beam temperature units (K km s^{-1}). This ratio is very similar that found by Saintonge et al. (2017) in their mass-selected sample of local galaxies and within the range found by Leroy et al. (2013) for nearby star-forming disc galaxies, indicating the molecular gas in NGC 0383 is similar to that in other local galaxies. The detection of CO line emission provides information about the cold gas mass distribution, that is later incorporated into our modelling (in addition to the kinematics themselves).
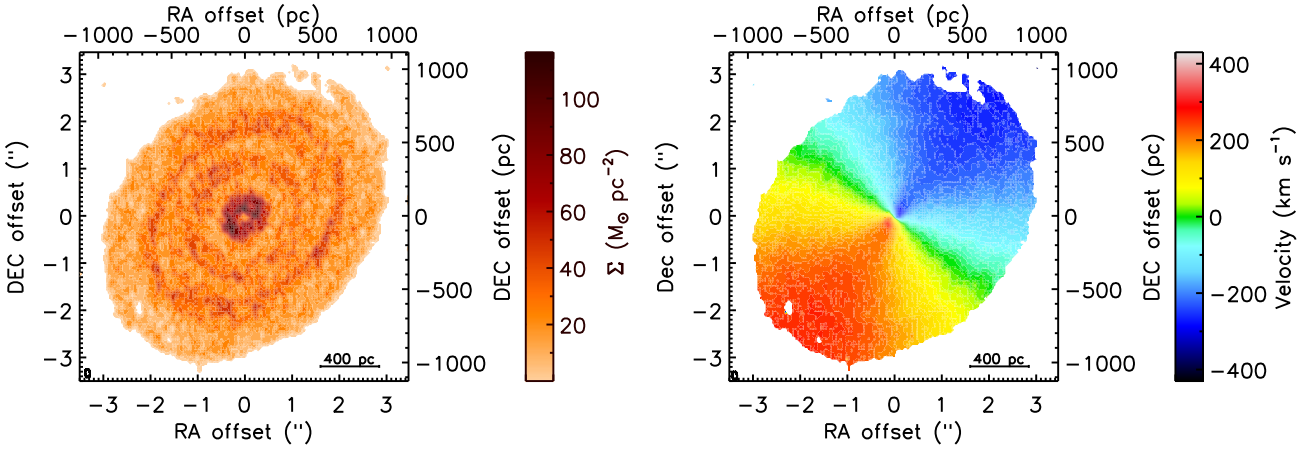
Figure 2. Moment zero (integrated intensity; left-hand panel) map of NGC 0383 assuming conversion factor CO-to-H$_2$ $\alpha_{CO} = 4.8\,M_\odot$(K km s^{-1})$^{-1}$ pc^{-2}. Moment one (intensity-weighted mean velocity; right-hand panel) of NGC 0383. The ellipse at the bottom-left of each panel shows the synthesised beam (0$\rlap{.}''$18×0$\rlap{.}''$1).
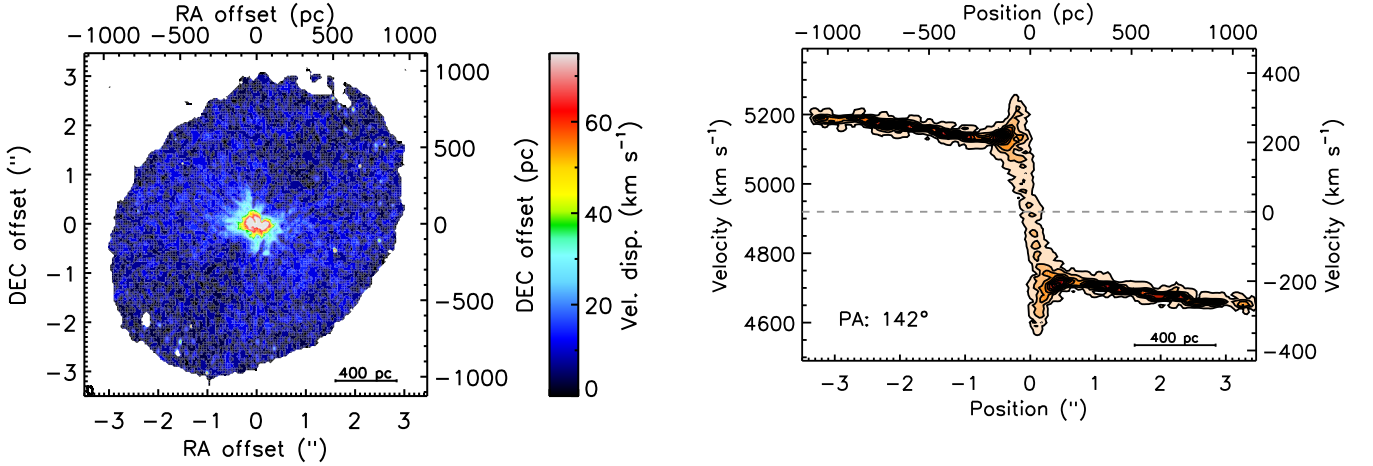

Figure 3. Moment two (intensity-weighted velocity dispersion) map of NGC 0383. The ellipse at the bottom-left shows the synthesised beam (0$\rlap{.}''$18×0$\rlap{.}''$1).

2.3 Continuum Emission

As mentioned previously, NGC 0383 hosts a radio-loud AGN. We detect a continuum point source at the kinematic centre of the galaxy, with a total integrated intensity of 65.2±0.1 mJy at a central frequency of 235.33 GHz. Adding to our flux those tabulated in the NASA/IPAC Extragalactic Database (NED)[2] at millimetre and radio wavelengths, we constructed a radio–sub-mm spectral energy distribution (SED), shown in Fig. 6. Our data point, shown by the cyan diamond, agrees well with previous observations. The literature data generally encompass emission from both the nucleus and the jets, but it is likely that it is the nucleus that causes the observed variability (i.e. the few data points well below the red best-fitting line in Fig. 6). Nevertheless, the data are fitted well with a simple power law for the flux F as a function of frequency ν ($F_\nu \propto \nu^\alpha$), with a power-law
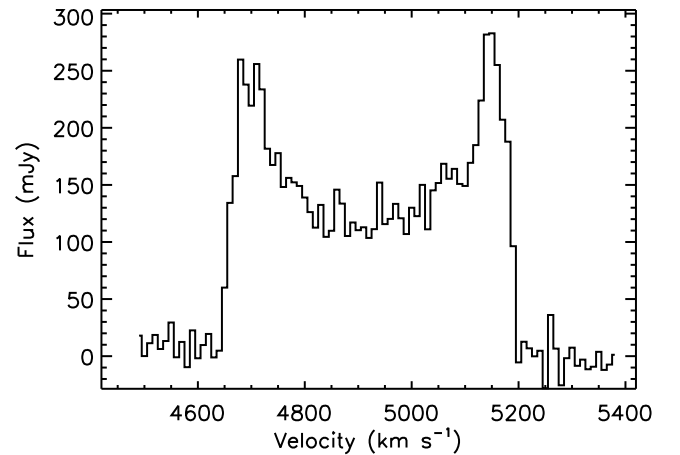

Figure 4. Major-axis position-velocity diagram of NGC 0383. The SMBH signature is clearly visible and dominant at radii less than 0$\rlap{.}''$5. The rotation of the outer disc ($\gtrsim$0$\rlap{.}''$5) is very regular and relaxed. The dashed line shows the systemic velocity $V = 4925 \pm 4$ km s^{-1}.

Figure 5. ^{12}CO(2–1) integrated spectrum of NGC 0383, showing the clear double-horned shape of a rotating disc.
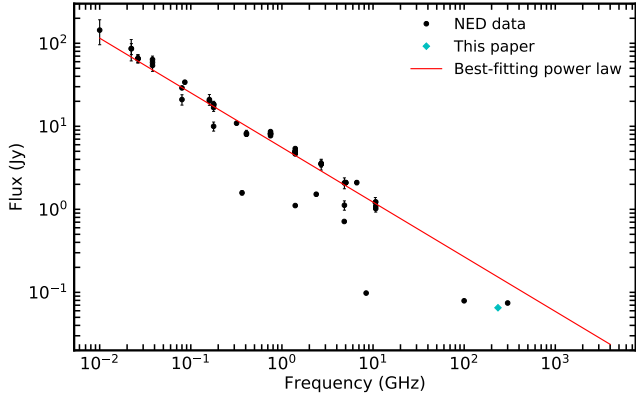
Figure 6. Spectral energy distribution of NGC 0383 from radio to mm wavelengths, constructed using data from NED (black circles) and our data (cyan diamond). A best-fitting power law with slope -0.66 is overlaid in red. Error bars are plotted for all points but most are smaller than the symbol used.

index $\alpha = -0.66 \pm 0.03$ (the red line shown in Fig. 6). This index value ($\approx$-0.7) is typical of a radio galaxy dominated by synchrotron radiation, as expected here from the prominent AGN jets (e.g. Macdonald et al. 1968; Bridle & Perley 1984; Laing & Bridle 2002).

Despite the prominence of the AGN jets, the extreme regularity of the molecular gas distribution and kinematics (Fig. 2) indicates that the radio AGN activity does not directly disturb the gas disc. Our ability to model the disc motions and estimate the SMBH mass is thus unaffected.

3   DYNAMICAL MODELLING

The method we use to estimate the SMBH mass is described in detail in Davis et al. (2017) and was used in the previous WISDOM papers, but we summarise the specifics for modelling NGC 0383 in this section. We make use of the publicly available KINEMATIC MOLECULAR SIMULATION (KinMS)[3] mm-wave observation simulation tool of Davis et al. (2013a) to create models of the data cube. KinMS uses input information about the gas distribution and kinematics, including a circular velocity curve. Applying observational effects such as beam smearing and velocity binning, KinMS then creates a simulated data cube that can be directly compared to the observed data cube. The model parameters are incrementally driven towards the best-fitting values by a Markov Chain Monte Carlo (MCMC) method. The MCMC algorithm fully samples the χ^2 hyper-volume to estimate the posterior distributions and hence uncertainties on the best-fitting values.

3.1   Mass Model

We use an axisymmetric model of the stellar light distribution to derive the circular velocity curve of the galaxy. We assume that the stellar mass dominates the potential in the inner parts of the galaxy: the molecular gas mass density is

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[3] https://github.com/TimothyADavis/KinMS

Table 1. MGE best-fitting Gaussians (not deconvolved).

I ($L_{\odot,\mathrm{F160W}}$ pc^{-2})	σ_j ($''$)	q_j
12913.05	0.0682	0.91
3996.80	0.823	0.9
5560.84	1.13	0.93
4962.29	2.63	0.9
2877.92	4.98	0.9
957.88	12.7	0.9

Notes: For each Gaussian component, column 1 lists its F160W central surface brightness, column 2 its standard deviation (width) and column 3 its axis ratio.

negligible in this system (see Section 4.2), while dark matter is usually unimportant at small radii, as shown by e.g. Cappellari et al. (2013). Even if this latter assumption is incorrect, if the dark matter were distributed identically to the stellar mass in the inner parts of the galaxy, it would simply lead to a higher mass-to-light ratio and would not affect the best-fitting SMBH mass. If dark matter were to contribute significantly and be distributed differently to the stellar mass, we would then find evidence for a significant mass-to-light ratio gradient (we find marginal evidence for a small mass-to-light ratio gradient in Section 3.2).

To model the luminous mass we perform a Multi-Gaussian Expansion (MGE; Emsellem et al. 1994), using the method implemented in the MGE_FIT_SECTORS Interactive Data Language (IDL) software[4] version v4.12 of Cappellari (2002). We use a combined *Hubble Space Telescope (HST)* Near Infrared Camera and Multi-Object Spectrometer (NICMOS) F160W and Two Micron All-Sky Survey (2MASS) *H*-band image. This combined image allows us to model the stellar light with a sum of two-dimensional (2D) Gaussians up to a radius of $20''$ (6.4 kpc), the *HST* image being used exclusively for the inner $\approx 4''$ (1.3 kpc) in radius because of its superior angular resolution. To minimise the effect of dust attenuation on the mass-to-light ratio, the *HST* image was masked over part of its lower-right limb (see the cyan region in Fig. 7, top panel). The resulting MGE model is shown in Fig. 7, with the values of each Gaussian listed in Table 1 (these values have not been deconvolved).

The circular velocity curve is then calculated by the MGE_CIRCULAR_VELOCITY procedure[4], by first analytically deprojecting the 2D Gaussians to a three-dimensional (3D) mass distribution, calculating the potential, and hence the circular velocity. The procedure uses a mass-to-light ratio of $1\,M_\odot/L_{\odot,\mathrm{F160W}}$. The circular velocity is then multiplied element-wise by the square root of the mass-to-light ratio adopted, and a point mass representing the SMBH is added in the centre. The functional form of the mass-to-light ratio is fully explained in Section 3.2. However, we will show in Section 4 that, in the case of NGC 0383, the SMBH mass is essentially independent of the stellar mass-to-light ratio.

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[4] http://purl.org/cappellari/software, part of the Jeans Anisotropic MGE (JAM) dynamical modelling package of Cappellari (2008).
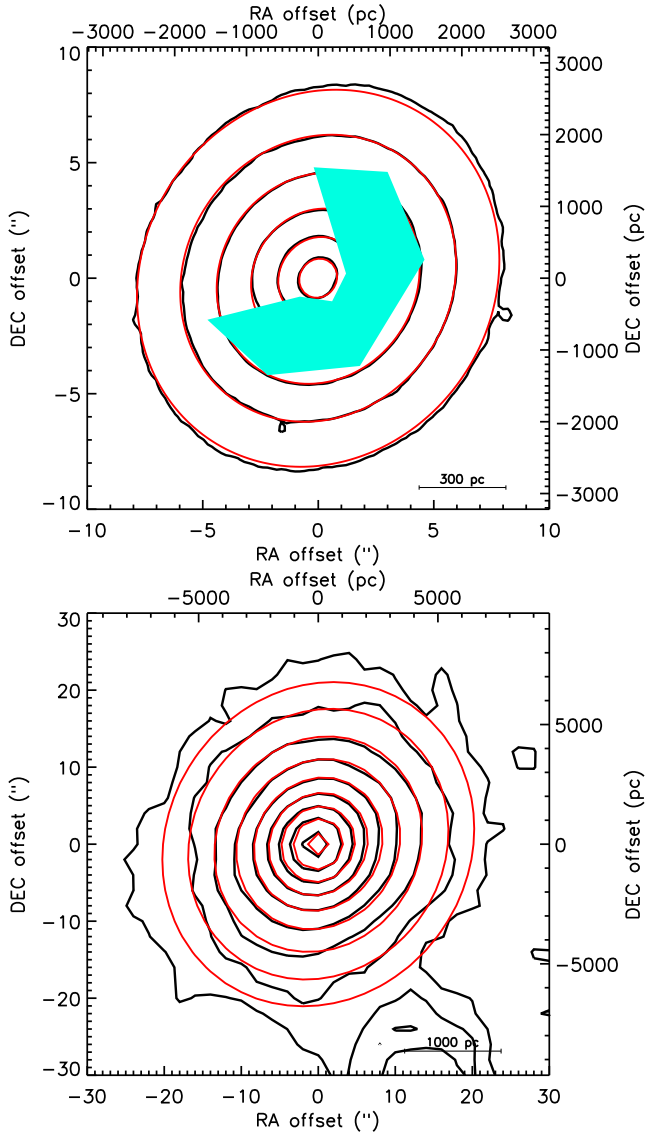
Figure 7. Multi-Gaussian Expansion (MGE) model of NGC 0383 (red contours) overlaid on the *HST* NICMOS F160W image (black contours, top panel) and the 2MASS *H*-band image (black contours, bottom panel). In the *HST* image (top panel), the area masked due to dust is shown in cyan. A foreground star, bottom-right in the 2MASS image (bottom panel), is outside the fit radius and does not affect the MGE.

3.2 Bayesian analysis

We use a MCMC method to find the posterior distribution of the model best fitting the NGC 0383 data, making use of the IDL KINMS_MCMC[5] code of Davis et al. (2013b, 2017), that easily interfaces with the KinMS simulation tool to create new models and calculate and maximise the likelihood. A single fit was made to the whole clean data cube, with one hundred thousand iterations. All parameters had flat priors in linear space within specified physical limits, as listed in Table 2, the only exception being the SMBH mass prior that

was flat in log-space. The observational errors were taken to be the RMS of the data cube in line-free channels, assumed to be constant throughout the cube.

The molecular gas disc of NGC 0383 has a slight nuclear ring and outer spiral/ring structures that make assuming a smoothly-varying monotonic radial profile inappropriate. Rather than construct an arbitrarily-complicated parametrisation of the radial gas distribution, we adopt instead the observed gas distribution as an input to our KINMS model. Using the SKYSAMPLER[6] tool (Smith et al. 2019), we thus sample the de-convolved CLEAN components produced by the CASA task to generate a set of gas particles that exactly replicate the surface brightness profile. These particles are then used as an input into KINMS, with the three-dimensional central position, inclination and position angle of the gas disc as free parameters. The centre is initially assumed to be at the centre of the continuum emission (RA= $01^h07^m24\overset{s}{.}95$, Dec.= $+32°24'45\overset{''}{.}15$) and the velocity of the central channel of the cube ($V_{\text{helio, radio}} = 4940$ km s^{-1}). With no evidence to the contrary, we use the thin disc approximation for NGC 0383.

We found that allowing a linearly-varying radial mass-to-light ratio profile fits the data better than a single (constant) mass-to-light ratio. Initial fits used a single mass-to-light ratio for the whole disc, but this did not provide a good fit to the entire data cube. We therefore implemented the simplest model to account for this, a linearly varying mass-to-light ratio, defined as

$$M/L(R) = (M/L_{\text{outer}} - M/L_{\text{inner}}) \left(\frac{R}{3\overset{''}{.}5} \right) + M/L_{\text{inner}}, \quad (2)$$

where R is the radius and the inner (M/L_{inner}) and outer (M/L_{outer}) mass-to-light ratios are free parameters of our fit. The inner value is set at the centre of the disc ($R = 0''$) with the outer edge at $R = 3\overset{''}{.}5$ and a flat mass-to-light ratio beyond that.

Here we adopt the usual definition of 1σ (3σ) uncertainties as the 68.3% (99.7%) confidence intervals of the Bayesian posteriors found from the MCMC. Table 2 lists the best-fitting value of each model parameter, along with its formal uncertainties.

As discussed in Section 3.2 of van den Bosch & van de Ven (2009), when working with very large datasets the statistical uncertainties can be severely underestimated due to the dominance of the systematic uncertainties. Accordingly, they suggest an approximate correction to account for the systematic uncertainties, by rescaling the $\Delta\chi^2$ (with respect to the minimum χ^2, χ^2_{min}) required to define a given confidence level by the standard deviation of the χ^2, namely $\sqrt{2(N-P)} \approx \sqrt{2N}$, where N is the number of constraints ($\approx 5.9 \times 10^6$) and P is the number of inferred model parameters (10). This sets the 68.3% (99.7%) confidence level at $\chi^2_{\text{min}} + \sqrt{2N}$ ($\chi^2_{\text{min}} + 3\sqrt{2N}$). Applying this rescaling results in significantly larger uncertainties on the fitted parameters, which are likely to be more physically plausible. The same method was applied and discussed in detail by Smith et al. (2019), and we use it here in the MCMC fitting of NGC 0383. The corner plots and one-dimensional marginalisation of each model parameter are shown in Fig. 8.

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[5] https://github.com/TimothyADavis/KinMS_MCMC

[6] https://github.com/Mark-D-Smith/KinMS-skySampler

Correlations are induced between pixels due to the synthesised beam, that can be corrected for by accounting for the induced covariance. However, the effect of this covariance on the MCMC uncertainties is negligible compared to the rescaling of the χ^2 discussed above, hence we did not include the covariance matrix in our calculations.

We find strong evidence for a SMBH, of mass $(4.2 \pm 0.7) \times 10^9 \, \mathrm{M_\odot}$ (3σ uncertainty). The best-fitting model's PVD is shown in the middle panel of Fig. 9 as the blue contours overlaid on the data. It has a reduced χ^2 of 1.01. Figure 9 shows that a kinematic model with a dark massive object at the centre is the only model to fully describe the data. In the left panel the SMBH has been removed and the model no longer reproduces the data. The right panel of Fig. 9 shows the best-fitting model with the mass-to-light ratio set to zero, i.e. no stellar mass, demonstrating the that SMBH mass dominates in the inner $0\rlap{.}''5$ (as the fit is still very good in that region).

The best-fitting F160W-band mass-to-light ratio decreases linearly from $2.8 \pm 0.6 \, \mathrm{M_\odot/L_{\odot, F160W}}$ in the centre to $2.4 \pm 0.3 \, \mathrm{M_\odot/L_{\odot, F160W}}$ at the outer edge of the disc (both 3σ uncertainties). The spatial centre (as indicated by the X and Y offsets) is consistent with the unresolved continuum source to within the beamsize.

4 DISCUSSION

In this work we have presented ALMA $^{12}\mathrm{CO}(2\text{–}1)$ observations of NGC 0383 showing a relaxed gas disc (Section 2). The data clearly show the kinematic signature of a massive dark object, with a mass of $(4.2 \pm 0.7) \times 10^9 \, \mathrm{M_\odot}$ (3σ uncertainty) measured through dynamical modelling (Section 3).

4.1 Uncertainties

The uncertainties associated with a SMBH mass derived through the molecular gas technique are discussed extensively in the previous papers using this method. Each paper builds from the last and focuses on the sources of uncertainty that are relevant to each galaxy. In particular, Smith et al. (2019) discussed properly constraining the mass-to-light ratio and inclination. As we resolve the SMBH R_{SOI} in NGC 0383 in this case our mass measurement is essentially independent of the mass-to-light ratio (see Section 4.2). Accurately determining the inclination, however, remains important. Our choice to apply the χ^2 scaling (discussed in Section 3.2) allows us to retrieve more physically meaningful estimates of our inclination uncertainties. As NGC 0383 is fairly face-on ($i \approx 38°$) these uncertainties dominate the error budget, through the degeneracy between inclination and SMBH mass (see Figure 8).

Other potential sources of uncertainty arise from the assumption that the molecular gas is dynamically cold and rotating on circular orbits. The velocity dispersion of the gas is consistently small ($\sigma_{\mathrm{gas}} < 10 \, \mathrm{km \, s^{-1}}$), indicating the disc is nearly perfectly rotationally supported ($V_{\mathrm{rot}}/\sigma_{\mathrm{gas}} \gtrsim 40$, where V_{rot} is the deprojected rotation velocity of the gas in the nearly flat portion of the rotation curve; see e.g. Fig. 4). Despite this, some non-circular motions do appear to be present. The velocity residuals (Fig. 10; data moment 1 minus best-fitting model moment 1) show the same spiral/ring structures noted in the moment zero (Fig. 2), indicating that material may be flowing along these arms (potential fuelling the AGN). The larger velocity residuals near the centre are due to the intensity weighting when creating the moment 1 map. However, the dominance of the SMBH in the central regions (see Fig. 11) indicates that non-circular motions are unlikely to significantly affect the derived SMBH mass.

All these uncertainties are small, and in any case they are dwarfed by that on the distance measurement. This uncertainty is $\approx 15\%$ (i.e. $66.6 \pm 9.9 \, \mathrm{Mpc}$), from the use of the Tully-Fisher relation in Freedman et al. (2001) to estimate the distance. The SMBH mass measurement scales linearly with the distance adopted, and as is customary we do not include the distance uncertainty in our result.

4.2 Mass-to-light Ratio Influence

The posterior distribution between SMBH mass and mass-to-light ratio shows a strong covariance (see Fig. 8, middle panels of the leftmost column), although this is exaggerated by plotting linear against logarithmic scales. The correlation present is contrary to the expected anti-correlation, and it may be a product of the SMBH mass–inclination and mass-to-light ratio–inclination correlations. By allowing the inclination to vary during the fit, the correlation between mass-to-light ratio and inclination dominates and induces correlations in other variables. See Smith et al. (2019) for a fuller discussion of this issue.

A simple calculation of the total mass enclosed from the circular velocity (and assuming spherical symmetry, i.e. $M_{<R} \propto V_{\mathrm{rot}}^2(R)/R$) allows us to determine how significant the stellar mass is as a function of radius in NGC 0383. Figure 11 shows the enclosed stellar mass as a function of radius as well as the enclosed total mass, revealing that the stellar mass becomes significant only at a radius of $\approx 300 \, \mathrm{pc}$. At $0\rlap{.}''13$ (i.e. one synthesised beam, $43 \, \mathrm{pc}$), the stellar mass is only $\approx 1\%$ of the mass enclosed at that radius, so is insignificant. The molecular gas disc mass at this radius is $\approx 10\%$ of the total enclosed mass, that is again small compared to the SMBH mass ($\approx 90\%$ of the total enclosed mass). This indicates that whilst the mass-to-light ratio (and assumed X_{CO}) do have a covariance with the SMBH, their effect on the best-fitting value is very small and the SMBH mass is largely independent of them. The fact that our SMBH mass is almost independent of our luminous mass model in turn leads to the very small uncertainties on M_{BH} compared to other works (indeed, the 3σ confidence interval in Fig. 8 is very narrow). It also gives us greater confidence in our measurement.

Although in NGC 0383 the SMBH dominates the total mass distribution within a few synthesised beams (i.e. angular resolution elements), this might not always be the case and in some instances the mass of the molecular gas disc itself may matter. In fact, even in NGC 0383, the stellar and molecular gas masses are approximately equal at R_{SOI}. This reinforces the importance of spatially-resolved molecular gas data for SMBH mass measurements.

4.3 Estimating M_{BH} from the Observed R_{SOI}

In contrast to the detailed dynamical modelling of Section 3.2, we can make a crude estimate of the SMBH mass

Table 2. Best-fitting parameters with uncertainties from the MCMC fits.

Parameter	Search range	Best fit	1σ uncertainty	3σ uncertainty
SMBH mass (log($M_\odot$))	8.70–9.95	9.63	0.04	0.08
Stellar M/L inner ($M_\odot/L_{\odot,\mathrm{F160W}}$)	0.01–10	2.78	0.21	0.61
Stellar M/L outer ($M_\odot/L_{\odot,\mathrm{F160W}}$)	0.01–10	2.36	0.12	0.33
Position angle (°)	112–172	142.20	0.04	0.10
Inclination (°)	26–89	37.58	1.67	3.48
Velocity dispersion (km s^{-1})	0–15	8.32	0.72	2.11
Nuisance parameters				
Integrated intensity (Jy km s^{-1})	5–200	74.60	4.15	9.79
Centre X offset ($''$)	-5–5	-0.00	0.01	0.03
Centre Y offset ($''$)	-5–5	-0.05	0.02	0.04
Centre velocity offset (km s^{-1})	-50–10	-15.16	1.37	3.63

Note: The X and Y offsets are measured with respect to the location of the unresolved continuum point source, RA= 01^{h}07^{m}24^{s}96 and Dec = +32°24′45″11. The velocity offset is measured with respect to the central channel of the cube ($V_{\mathrm{helio,radio}} = 4940$ km s^{-1}). The best-fitting centre velocity offset thus defines a systemic velocity of $V = 4925 \pm 4$ km s^{-1}.

from the observed radius of the SMBH sphere of influence (R_{SOI}). The SMBH R_{SOI} is defined as the radius within which the SMBH dominates the potential (see Eq. 1). This radius can be determined from the observed PVD (Fig. 4) as the local minimum in the rotation curve (i.e the transition point) between the SMBH-dominated Keplerian curve ($V_{\mathrm{rot}} \propto 1/\sqrt{R}$) and the stellar mass-dominated approximately flat rotation curve ($V_{\mathrm{rot}} \simeq$ constant). By visual inspection, we estimate this occurs at a radius of $\approx 0''7$ (see Fig. 9). Using Eq. 1 and $\sigma_* = 239 \pm 16$ km s^{-1} (van den Bosch 2016) then yields $M_{\mathrm{BH}} \approx 3.0 \times 10^9$ M$_\odot$. Given that the SMBH R_{SOI} is so well resolved, this back-of-the-envelope estimate agrees well with our full modelling value of $(4.2 \pm 0.7) \times 10^9$ M$_\odot$ (the latter also yielding an exact $R_{\mathrm{SOI}} = 316 \pm 60$ pc or $0''98 \pm 0''18$).

4.4 Comparison to the Literature

An upper limit on the SMBH mass in NGC 0383 has previously been determined by Beifiori et al. (2009). Once scaled to our distance (66.6 Mpc from 63.4 Mpc) and inclination (37°6 from 33°) this is $M_{\mathrm{BH}} = 1.1 \times 10^9$ M$_\odot$. Given that this is lower than our measurement, it might indicate the presence of very disturbed ionised gas.

One of the tightest known correlations between SMBH mass and a host galaxy property is that with the stellar velocity dispersion, i.e. the $M_{\mathrm{BH}} - \sigma_*$ relation (e.g. Gebhardt et al. 2000; Ferrarese & Merritt 2000). We added our measurement to the dynamical measurements and power-law fit of van den Bosch (2016) in Fig. 12, to see whether our measurement also lies on this relation. The data of van den Bosch (2016) are shown in grey, while our new measurement for NGC 0383 is shown in blue. Other SMBH masses estimated using the molecular gas method are shown in red (Davis et al. 2013b; Onishi et al. 2015; Barth et al. 2016a,b; Onishi et al. 2017; Davis et al. 2017, 2018; Smith et al. 2019; Combes et al. 2019; Nagai et al. 2019; Boizelle et al. 2019). NGC 0383 has the largest SMBH mass estimated with molecular gas so far, and is on the upper edge of the scatter in the van den Bosch (2016) M_{BH}-σ_* relation. If accurate, the low σ_e compared to the SMBH mass indicates NGC 0383 might be part of the so-called "over-massive black hole" population. These galaxies are thought to be local analogues of the higher-redshift quiescent galaxies that also contain over-massive

black holes, and could therefore be relics that have evolved little since a redshift $z \approx 2$ (van den Bosch 2016; Walsh et al. 2016). Alternatively, an over-massive SMBH may be the result of merger(s), especially when paired with a high molecular gas mass in an ETG harbouring an AGN. Lim et al. (2000) discuss whether the AGN activity in NGC 0383 is the result of a gas-rich minor merger, with black hole accretion beginning soon after the merger event. While this is plausable, it is unclear if this scenario could explain such an overly massive SMBH.

4.5 Comparison of Spatial Scales Probed by Molecular Gas and Megamasers

Modelling megamaser dynamics is typically the most accurate method of measuring SMBH masses, due to the exquisite angular and spectral resolution usually achieved. The Keplerian rise we detect in the centre of NGC 0383 indicates that our data reach very close to the SMBH. The connection between the accretion disc/torus region (where masers are typically found) and the outer molecular gas disc has only recently begun to be explored in any detail. It is thought that position angle mismatches are common between these two components, and tilted and counter-rotating accretion discs are frequently observed (e.g. recently by Imanishi et al. 2018; Combes et al. 2019). In contrast to this expectation, NGC 0383 seems to have a single, unwarped molecular disc extending from kiloparsec scale to well within its SMBH SOI.

We are able to estimate just how close to the SMBH the highest velocity molecular gas we detect here is. Equating the centrifugal and gravitational forces at a radius R and assuming the SMBH mass dominates the stellar mass within this radius, we obtain

$$R = \frac{GM_{\mathrm{BH}}}{V_c^2}, \tag{3}$$

where V_c is the circular velocity at R. If we normalise the radius by the Schwarzschild radius

$$R_{\mathrm{Schw}} \equiv \frac{2GM_{\mathrm{BH}}}{c^2}, \tag{4}$$

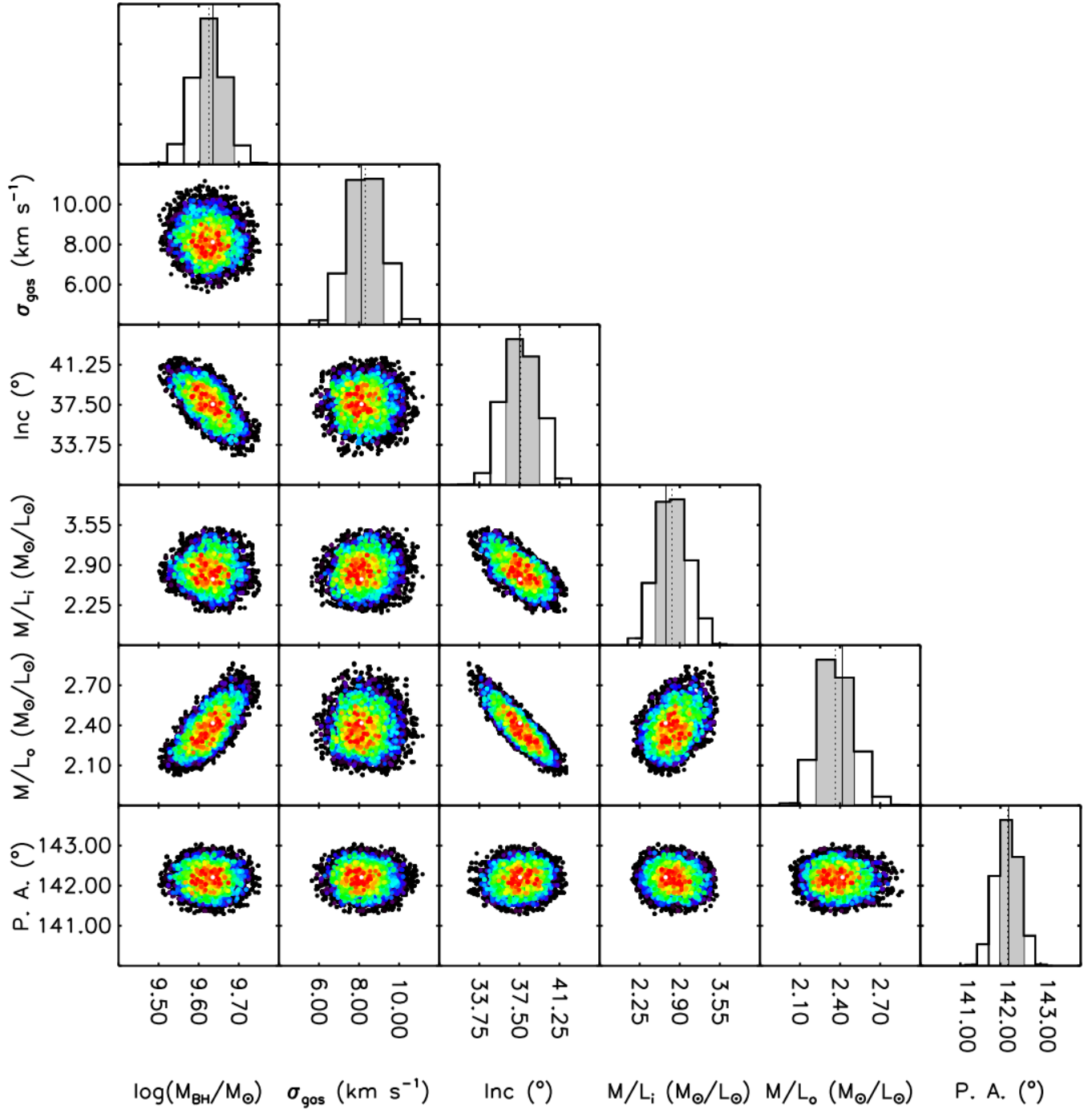where c is the speed of light, and the circular velocity by c,

Figure 8. Corner plots showing covariances between the model parameters, for the non-nuisance parameters. The colours represent increasing confidence intervals from 68.3% (red, 1σ) to 99.7% (blue, 3σ). The white dots show the χ^2_{min} values. Covariances are present between the SMBH mass and outer stellar mass-to-light ratio, SMBH mass and inclination, and inclination and both stellar mass-to-light ratios. In the SMBH mass cases, this is exaggerated by plotting linear against logarithmic scales. Histograms show the one-dimensional marginalised posterior distribution of each model parameter. The shaded regions indicate the 68% (1σ) confidence intervals. The black dashed lines show the median values and the black solid lines the χ^2_{min} values.

we find

$$\frac{R}{R_{\rm Schw}} = \frac{1}{2}\left(\frac{V_{\rm c}}{c}\right)^{-2}, \tag{5}$$

where interestingly the SMBH mass has dropped out. All rotationally-supported discs around a SMBH should thus follow this unique relation, irrespective of the SMBH mass.

Substituting $V_{\rm c} = V_{\rm obs}/\sin(i)$, where $V_{\rm obs}$ is the observed line-of-sight velocity (along the galaxy major axis) and i the inclination, we obtain

$$\frac{R}{R_{\rm Schw}} = 0.5\times 10^6 \left(\frac{300\,{\rm km\,s^{-1}}\sin(i)}{V_{\rm obs}}\right)^2. \tag{6}$$

The maximum rotation velocity observed in NGC 0383
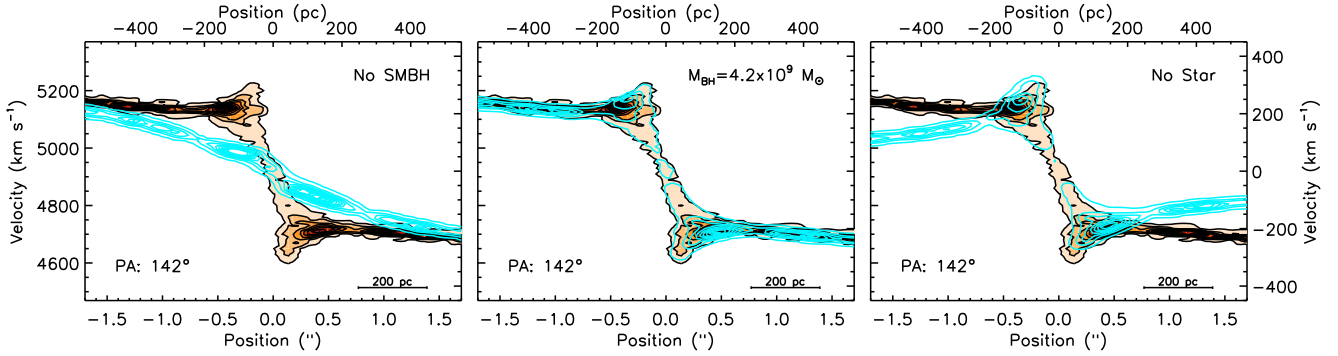
Figure 9. Observed position-velocity diagram (PVD) of NGC 0383 with the smooth mask applied (orange contours) with the best-fitting model's PVD overplotted (blue contours; middle panel). The left panel shows the same model, but with no SMBH. The right panel also shows the same model, but with both mass-to-light ratios set to zero (no stellar contribution).
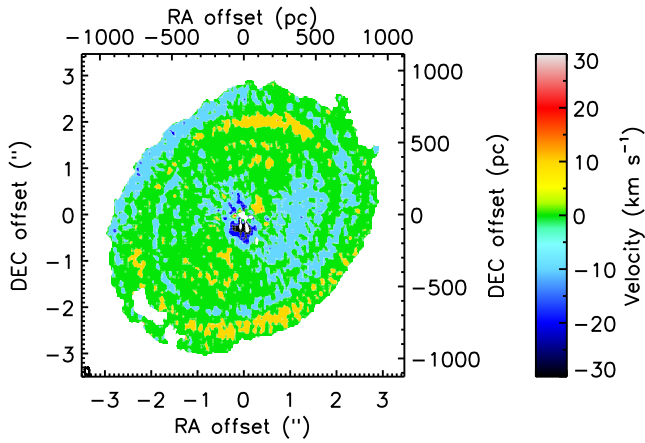
Figure 10. First moment (intensity-weighted mean velocity) residuals of NGC 0383, created by subtracting the first moment of the best-fitting model cube from the first moment of the data cube. The plot clearly shows the slight spiral features that could not be modelled by our axisymmetric mass model. Due to the simplicity of the model and weighting when creating the first moment, there are larger velocity residuals near the disc centre.

is $V_{\rm obs} \approx 350\,{\rm km\,s^{-1}}$ (the peak of the PVD in Fig. 4) and $i = 37.^{\circ}5$. The highest velocity molecular gas we detect therefore reaches $\approx 1.36 \times 10^5$ Schwarzschild radii.

Megamasers, although rare, are the current gold standard for dynamical SMBH mass measurements. Megamasers are thought to trace gas very close to the SMBH (in the accretion disc/torus), and as such they probe the gravitational field of the SMBH in a way that is unaffected by most outside sources. In addition, in the best cases, maser observations provide independent geometric distance estimates, vastly reducing the dominant systematic effect that plagues most SMBH mass measurements. Some of the earliest megamasers discovered were in NGC 4258 (Nakai et al. 1993; Herrnstein et al. 1999), more recently The Megamaser Cosmology Project (MCP) have carried out the most complete survey of megamasers to date, with the goal of measuring Hubble's constant (see e.g. the survey compilation by Braatz et al. 2015). The MCP observations also allow them to make several SMBH mass measurements (e.g. Reid et al. 2009; Zhao et al. 2018). The observed megamasers with SMBH masses have $V_{\rm obs}$ ranging from $170\,{\rm km\,s^{-1}}$ (NGC 1029; Gao
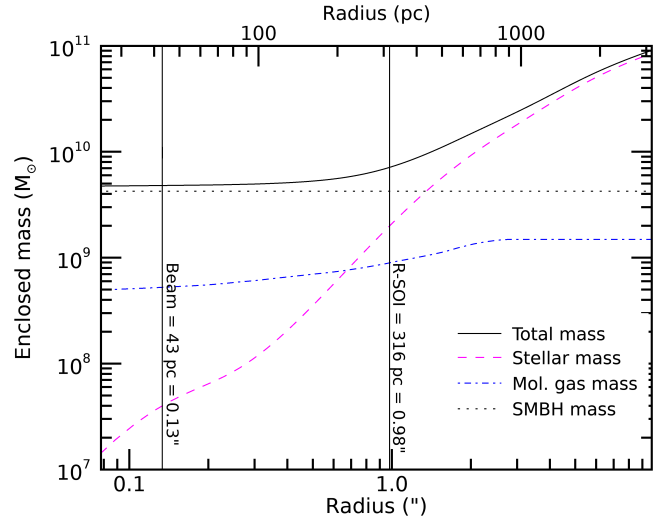
Figure 11. Cumulative mass profile of NGC 0383, showing the total mass (black solid line), SMBH mass (black dotted line), stellar mass (magenta dashed line) and molecular gas disc mass (blue dot-dashed line) enclosed as a function of radius. The vertical lines indicate the synthesised beam and measured $R_{\rm SOI}$. At a radius of one synthesised beam, both the stellar mass and the molecular gas mass are insignificant ($\approx 1\%$ and $\approx 10\%$ of the total mass, respectively).

et al. 2017) to $950\,{\rm km\,s^{-1}}$ (NGC 2273; Kuo et al. 2011), with an average of $\approx 600\,{\rm km\,s^{-1}}$. All megamaser systems are observed close to edge-on, so that $\sin(i) \approx 1$. Given this, megamasers typically probe gas at radii between 5×10^4 and 1.5×10^6 Schwarzschild radii. Our data thus show that the molecular gas disc in NGC 0383 extends unbroken and unwarped down to very close to the SMBH, and that it traces the same material probed by megamasers in other galaxies.

5 CONCLUSIONS

We have presented a measurement of the mass of the SMBH in the nearby lenticular galaxy NGC 0383 (radio source 3C 031). This estimate is based on ALMA observations of the ^{12}CO(2–1) emission line with a physical resolution of $\approx 43\,{\rm pc}$ ($0.''18 \times 0.''1$). We thus have a spatial resolution a
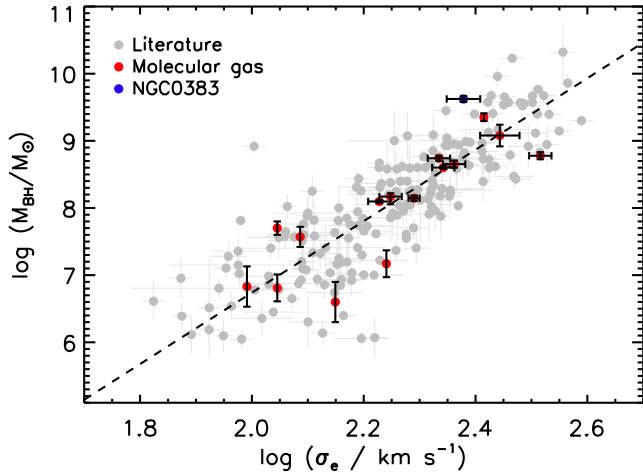
Figure 12. The $M_{BH} - \sigma_e$ relation from literature measurements (grey points and dotted line), as compiled by van den Bosch (2016). SMBH mass measurements using the molecular gas method are highlighted in red, while this work (NGC 0383) is in blue. For the molecular gas-derived SMBH masses, the error bars shown correspond to 1σ uncertainties.

factor of > 7 better than the R_{SOI}. Our spectroscopic resolution, and a channel width of $10\,km\,s^{-1}$, allow us to resolve gas down to $\approx 140,000$ Schwarzschild radii and thus to probe the same material as typical megamaser observations. NGC 0383 has a relaxed, smooth nuclear disc with weak ring/spiral features. We detect a clear Keplerian increase of the rotation velocity of $^{12}CO(2-1)$ at radii $\lesssim 0''.5$, and forward model of our ALMA data cube with the KinMS tool in a Bayesian MCMC framework to measure a SMBH mass of $(4.2 \pm 0.7) \times 10^9\,M_\odot$, a F160W-band mass-to-light ratio varying linearly from $2.8 \pm 0.6\,M_\odot/L_{\odot,F160W}$ in the centre to $2.4 \pm 0.3\,M_\odot/L_{\odot,F160W}$ at the outer edge of the molecular gas disc ($3''.5$ radius) and a velocity dispersion of $8.3 \pm 2\,km\,s^{-1}$ (all 3σ uncertainties). We also detect continuum emission from the AGN in NGC 0383 across the full bandwidth, consistent with synchrotron radiation. This work not only shows the power of ALMA to estimate SMBH masses, but it also demonstrates that the molecular gas method is highly complimentary to megamaser observations as it can probe the same emitting material.

REFERENCES

Alatalo K., et al., 2013, MNRAS, 432, 1796
Barth A. J., Boizelle B. D., Darling J., Baker A. J., Buote D. A., Ho L. C., Walsh J. L., 2016a, ApJ, 822, L28
Barth A. J., Darling J., Baker A. J., Boizelle B. D., Buote D. A., Ho L. C., Walsh J. L., 2016b, ApJ, 823, 51
Beifiori A., Sarzi M., Corsini E. M., Bontà E. D., Pizzella A., Coccato L., Bertola F., 2009, ApJ, 692, 856
Beifiori A., Courteau S., Corsini E. M., Zhu Y., 2012, MNRAS, 419, 2497
Bennett A. S., 1962, Mem. RAS, 68, 163
Boizelle B. D., Barth A. J., Walsh J. L., Buote D. A., Baker A. J., Darling J., Ho L. C., 2019, ApJ, 881, 10
Braatz J., et al., 2015, IAU General Assembly, 22, 2255730
Bridle A. H., Perley R. A., 1984, ARA&A, 22, 319
Bundy K., et al., 2008, ApJ, 681, 931
Cappellari M., 2002, MNRAS, 333, 400
Cappellari M., 2008, MNRAS, 390, 71
Cappellari M., 2016, ARA&A, 54, 597
Cappellari M., Verolme E. K., van der Marel R. P., Kleijn G. A. V., Illingworth G. D., Franx M., Carollo C. M., de Zeeuw P. T., 2002, ApJ, 578, 787
Cappellari M., et al., 2013, MNRAS, 432, 1709
Combes F., et al., 2019, A&A, 623, A79
Courteau S., et al., 2014, Reviews of Modern Physics, 86, 47
Dame T. M., 2011, preprint (arXiv:1101.1499),
Davis T. A., 2014, MNRAS, 443, 911
Davis T. A., et al., 2013a, MNRAS, 429, 534
Davis T. A., Bureau M., Cappellari M., Sarzi M., Blitz L., 2013b, Nature, 494
Davis T. A., Bureau M., Onishi K., Cappellari M., Iguchi S., Sarzi M., 2017, MNRAS, 468, 4675
Davis T. A., et al., 2018, MNRAS, 473, 3818
Djorgovski S., Davis M., 1987, ApJ, 313, 59
Dressler A., Richstone D. O., 1988, ApJ, 324, 701
Dressler A., Lynden-Bell D., Burstein D., Davies R. L., Faber S. M., Terlevich R., Wegner G., 1987, ApJ, 313, 42

Edge D. O., Shakeshaft J. R., McAdam W. B., Baldwin J. E., Archer S., 1959, Mem. RAS, 68, 37

Emsellem E., Monnet G., Bacon R., 1994, A&A, 285, 723

Ferrarese L., Merritt D., 2000, ApJ, 539, L9

Ferrarese L., Ford H. C., Jaffe W., 1996, ApJ, 470, 444

Freedman W. L., et al., 2001, ApJ, 553, 47

Gao F., et al., 2017, ApJ, 834, 52

Gebhardt K., et al., 2000, ApJ, 539, L13

Graham A. W., 2007, MNRAS, 379, 711

Greene J. E., et al., 2010, ApJ, 721, 26

Haering N., Rix H.-W., 2004, ApJ, 604, L89

Healey S. E., Romani R. W., Taylor G. B., Sadler E. M., Ricci R., Murphy T., Ulvestad J. S., Winn J. N., 2007, ApJS, 171, 61

Herrnstein J. R., et al., 1999, Nature, 400, 539

Högbom J. A., 1969, in Bulletin of the American Astronomical Society. p. 348

Hudson M. J., Lucey J. R., Smith R. J., Schlegel D. J., Davies R. L., 2001, MNRAS, 327, 265

Imanishi M., Nakanishi K., Izumi T., Wada K., 2018, ApJ, 853, L25

Jorgensen I., Franx M., Kjaergaard P., 1995, MNRAS, 276, 1341

Kormendy J., Ho L. C., 2013, ARA&A, 51, 511

Kormendy J., Richstone D., 1995, ARA&A, 33, 581

Krajnović D., McDermid R. M., Cappellari M., Davies R. L., 2009, MNRAS, 399, 1839

Krajnović D., Cappellari M., McDermid R. M., 2018, MNRAS, 473, 5237

Kuo C. Y., et al., 2011, ApJ, 727, 20

Laing R. A., Bridle A. H., 2002, MNRAS, 336, 1161

Leroy A. K., et al., 2013, AJ, 146, 19

Lim J., Leon S., Combes F., Dinh-V-Trung 2000, ApJ, 545, L93

Lo K. Y., 2005, ARA&A, 43, 625

Macdonald G. H., Kenderdine S., Neville A. C., 1968, MNRAS, 138, 259

Madau P., Dickinson M., 2014, ARA&A, 52, 415

Magorrian J., et al., 1998, AJ, 115, 2285

Marconi A., Hunt L. K., 2003, ApJ, 589, L21

McConnell N. J., Ma C.-P., 2013, ApJ, 764

McMullin J. P., Waters B., Schiebel D., Young W., Golap K., 2007, in Shaw R. A., Hill F., Bell D. J., eds, Astronomical Society of the Pacific Conference Series Vol. 376, Astronomical Data Analysis Software and Systems XVI. p. 127

Miyoshi M., Moran J., Herrnstein J., Greenhill L., Nakai N., Diamond P., Inoue M., 1995, Nature, 373, 127

Mullaney J. R., et al., 2012, ApJ, 753, 30

Nagai H., et al., 2019, preprint (arXiv:1905.06017),

Nakai N., Inoue M., Miyoshi M., 1993, Nature, 361, 45

Ocaña Flaquer B., Leon S., Combes F., Lim J., 2010, A&A, 518, A9

Okuda T., Kohno K., Iguchi S., Nakanishi K., 2005, ApJ, 620, 673

Onishi K., Iguchi S., Sheth K., Kohno K., 2015, ApJ, 806, 39

Onishi K., Iguchi S., Davis T. A., Bureau M., Cappellari M., Sarzi M., Blitz L., 2017, MNRAS, 468, 4663

Regan M. W., Thornley M. D., Helfer T. T., Sheth K., Wong T., Vogel S. N., Blitz L., Bock D. C. J., 2001, ApJ, 561, 218

Reid M. J., Braatz J. A., Condon J. J., Greenhill L. J., Henkel C., Lo K. Y., 2009, ApJ, 695, 287

Saintonge A., et al., 2017, ApJS, 233, 22

Sarzi M., Rix H., Shields J. C., Rudnick G., Ho L. C., McIntosh D. H., Filippenko A. V., Sargent W. L. W., 2001, ApJ, 550, 65

Scharwächter J., McGregor P. J., Dopita M. A., Beck T. L., 2013, MNRAS, 429, 2315

Smith M. D., et al., 2019, MNRAS, 485, 4359

Tully R. B., Fisher J. R., 1977, A&A, 54, 661

Walsh J. L., Barth A. J., Ho L. C., Sarzi M., 2013, ApJ, 770, 86

Walsh J. L., van den Bosch R. C. E., Gebhardt K., Yıldırım A., Richstone D. O., Gültekin K., Husemann B., 2016, ApJ, 817, 2

Zhao W., et al., 2018, ApJ, 854, 124

van Velzen S., Falcke H., Schellart P., Nierstenhöfer N., Kampert K.-H., 2012, A&A, 544, A18

van den Bosch R. C. E., 2016, ApJ, 831, 134

van den Bosch R. C. E., van de Ven G., 2009, MNRAS, 398, 1117

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