



# Erratum: “The Evolution of NGC 7465 as Revealed by Its Molecular Gas Properties” (2021, ApJ, 909, 98)

Lisa M. Young<sup>1,2</sup> , David S. Meier<sup>1,2</sup> , Martin Bureau<sup>3</sup> , Alison Crocker<sup>4</sup> , Timothy A. Davis<sup>5</sup> , and Selçuk Topal<sup>6</sup> <sup>1</sup> Physics Department, New Mexico Institute of Mining and Technology, 801 Leroy Place, Socorro, NM 87801, USA; [lisa.young@nmt.edu](mailto:lisa.young@nmt.edu)<sup>2</sup> Adjunct Astronomer, National Radio Astronomy Observatory, Socorro, NM 87801, USA<sup>3</sup> Sub-department of Astrophysics, Department of Physics, University of Oxford, Denys Wilkinson Building, Keble Road, Oxford, OX1 3RH, UK<sup>4</sup> Department of Physics, Reed College, Portland, OR 97202, USA<sup>5</sup> School of Physics & Astronomy, Cardiff University, Queens Buildings, The Parade, Cardiff, CF24 3AA, UK<sup>6</sup> Department of Physics, Van Yüzüncü Yıl University, Van 65080, Turkey

Received 2022 September 6; published 2022 September 26

In Table 3 of the published article we presented calculations for the dynamical and molecular masses of the unidentified spectral line source behind NGC 7465. However, we made errors in the calculation of the line width and the line flux, and those errors propagated through to the dynamical masses, line luminosities, molecular masses, and molecular/dynamical mass ratios. Hogg (1999) and Bourne et al. (2019) give brief explanations of some issues involved in estimating line widths at high redshift. Following their methods, the line width of the unidentified object (0.14 GHz at an observed frequency of 97.67 GHz) corresponds to a velocity width in the object’s rest frame of  $430 \text{ km s}^{-1}$ . Its integrated line flux is then  $0.89 \pm 0.09 \text{ Jy km s}^{-1}$ . We present updated values for the other parameters in Table 3. Because the molecular-to-dynamical mass ratios are smaller than in the previous version, the plausibility of the source having a redshift  $z > 1$  is larger, but none of the conclusions of the published article are changed.

**Table 3**  
Flux, Luminosity, and Mass Estimates for the Unidentified Source

Line ID	$z$	Diameter (kpc)	$M_{\text{dyn}}$ ( $M_{\odot}$ )	$L'_{\text{line}1}$ ( $\text{K km s}^{-1} \text{ pc}^2$ )	$M_{\text{mol}}$ ( $M_{\odot}$ )	$M_{\text{mol}}/M_{\text{dyn}}$
$^{12}\text{CO} (1-0)$	0.1802	1.06	$5.7 \times 10^9 / \sin^2 i$	$1.4 \times 10^9$	$5.1 \times 10^9$	$0.89 \sin^2 i$
$^{12}\text{CO} (2-1)$	1.3604	2.94	$1.6 \times 10^9 / \sin^2 i$	$2.2 \times 10^{10}$	$1.0 \times 10^{11}$	$6.5 \sin^2 i$
$^{12}\text{CO} (3-2)$	2.5404	2.81	$1.5 \times 10^9 / \sin^2 i$	$3.0 \times 10^{10}$	$1.1 \times 10^{11}$	$7.0 \sin^2 i$

**Note.** Calculations are made assuming a flat universe with  $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ ,  $\Omega_m = 0.3$ , and  $\Omega_{\Lambda} = 0.7$ . The linear diameter is estimated from the emission centroids in the outermost channels. The line luminosity  $L'_{\text{line}}$  is calculated as in Carilli & Walter (2013). Estimated luminosity conversions from  $J = 2-1$  or  $J = 3-2$  to  $J = 1-0$  are made using assumed excitations as in Boogaard et al. (2019), and the molecular mass (with He) is then estimated from the inferred  $^{12}\text{CO} (1-0)$  luminosity using a conversion factor  $\alpha = 3.6 M_{\odot} (\text{K km s}^{-1} \text{ pc}^2)^{-1}$  (Boogaard et al. 2019).



Original content from this work may be used under the terms of the [Creative Commons Attribution 4.0 licence](https://creativecommons.org/licenses/by/4.0/). Any further distribution of this work must maintain attribution to the author(s) and the title of the work, journal citation and DOI.

### ORCID iDs

Lisa M. Young  <https://orcid.org/0000-0002-5669-5038>  
David S. Meier  <https://orcid.org/0000-0001-9436-9471>  
Martin Bureau  <https://orcid.org/0000-0003-4980-1012>  
Alison Crocker  <https://orcid.org/0000-0001-8513-4945>  
Timothy A. Davis  <https://orcid.org/0000-0003-4932-9379>  
Selçuk Topal  <https://orcid.org/0000-0003-2132-5632>

### References

- Boogaard, L. A., Decarli, R., González-López, J., et al. 2019, *ApJ*, **882**, 140  
Bourne, N., Dunlop, J. S., Simpson, J. M., et al. 2019, *MNRAS*, **482**, 3135  
Carilli, C. L., & Walter, F. 2013, *ARA&A*, **51**, 105  
Hogg, D. W. 1999, arXiv:[astro-ph/9905116](https://arxiv.org/abs/astro-ph/9905116)